

at Observatorio do Pico dos Dias MCT/LNA, Brazil. The object shows clear signatures of tidal perturbations, i.e., two rings (one of them resembling a perturbed arm-like structure), a large visible envelope surrounding the whole galaxy, and a smooth tidal filament to the SE. Some smaller structures have also been highlighted using the GMOS images. A second nucleus-like object is visible on the SW of the ESO 287-G40 nucleus. The B-R versus B-V diagnostic diagram shows that the NE section of the external ring is populated with well-defined H II regions (active star-forming clumps) and is bluer and clumpier than the SW region, which is redder and smoother, without clear evidences of H II regions. Both rings are evidences of dramatic double triggering of starbursts due to density wave shocks throughout the galactic disk. The internal ring is redder than the external one, and the nucleus is redder than the secondary nucleus-like object. The smooth tidal filament to the SE shows evidences of young and old populations mixture, suggesting star and gas whirlwinds triggered by tidal interaction. We present the results of surface photometry, galaxy mass calculations and estimated dimensions of the highlighted structures.

PAINEL 150

STELLAR POPULATIONS IN THE KINEMATICAL SUBSYSTEMS OF TWO NEARBY ELLIPTICALS

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The distribution of the stellar populations in a spheroidal system and their kinematical properties supply important informations about the formation process. The spatial distributions of the mean luminosity-weighted stellar age, metallicity, and α/Fe ratio along both photometric axes of two nearby elliptical galaxies have been obtained using Lick index measurements on long slit spectra in order to reconstruct the star formation history in their kinematically distinct subsystems. Lick indexes were compared with those of single-aged stellar population (SSP) models. A population synthesis method was also applied in order to help disentangling the age-metallicity degeneracy of the Lick indexes. The star characteristics are associated with their kinematics: they are older and α -enhanced in the not rotating bulge of NGC 1052 and counter rotating core of NGC 7796, while they show a strong spread of α/Fe and age along the rotating disk of NGC 1052 and an outwards radial decreasing of them outside the core of NGC 7796. Both galaxies were formed by processes in which the star formation occurred firstly at the bulge (NGC 1052) and nucleus (NGC 7796) 12-15 Gyr ago on short timescales (0.1-1 Gyr) providing an efficient chemical enrichment by SN-II, differently of the other kinematical regions. (A. Milone acknowledges FAPESP, 2006/05029-3 and 2000/06695-0, and *Proap-Capes/INPE*).

PAINEL 151

MASS DISTRIBUTION IN HICKSON COMPACT GROUPS

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We are presenting in this poster the results of the last part of our study on the influence on dense environment on the kinematics and Dynamics of Hickson Compact Groups of Galaxies. Mendes de Oliveira et al. (2003) showed in the study that galaxies in CG follow the B band T-F relation as field galaxies, indicating the these galaxies show common mass - size relations and that the halo of HCG galaxies have been Significantly stripped inside R25. In order complete this work and understand better the mass distribution of galaxies inside dense group, we decide to analyse the mass distribution using the Rotation Curves combined with the J band photometry in order to get the older stars. We used the model desolved by Blais-Ouellette et al. (1999) where the photometric profile is decomposed into bulge, disk and spherical halo and the photometric parameters used to fit the observed Rotation Curves. We modeled 20 of the 25 galaxies of our sample and we are presenting the correlations between the different parameters for the halo, the bulge and the disk, such as, the central halo density, the central disk surface brightness, the different mass to light ratios and characteristics radius.

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ESPECTROSCÓPIA GMOS-SUL DA GALÁXIA ANELADA PECULIAR FM 287-14

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Galáxias aneladas peculiares (GAp) possuem, geralmente, um núcleo fora do centro geométrico do anel e representam locais de vigorosa formação estelar não nuclear. Em geral, esses anéis apresentam nódulos (regiões de intensa formação