

## GRB 140619B: A SHORT GRB FROM A BINARY NEUTRON STAR MERGER LEADING TO BLACK HOLE FORMATION

R. RUFFINI<sup>1,2,3,4</sup>, M. MUCCINO<sup>1,2</sup>, M. KOVACEVIC<sup>1,3</sup>, F. G. OLIVEIRA<sup>1,3</sup>, J. A. RUEDA<sup>1,2,4</sup>, C. L. BIANCO<sup>1,2</sup>,

M. ENDERLI<sup>1,3</sup>, A. V. PENACCHIONI<sup>4,5</sup>, G. B. PISANI<sup>1,2</sup>, Y. WANG<sup>1,2</sup>, AND E. ZANINONI<sup>4</sup>

<sup>1</sup> Dip. di Fisica and ICRA, Sapienza Università di Roma, Piazzale Aldo Moro 5, I-00185 Rome, Italy; ruffini@icra.it

<sup>2</sup> ICRA Net, Piazza della Repubblica 10, I-65122 Pescara, Italy

<sup>3</sup> Université de Nice Sophia Antipolis, CEDEX 2, Grand Château Parc Valrose, Nice, France

<sup>4</sup> ICRA Net-Rio, Centro Brasileiro de Pesquisas Físicas, Rua Dr. Xavier Sigaud 150, Rio de Janeiro, RJ, 22290-180, Brazil

<sup>5</sup> Instituto Nacional de Pesquisas Espaciais, Av. dos Astronautas, 1758, São José dos Campos, SP, 12227-010, Brazil

Received 2014 November 30; accepted 2015 June 15; published 2015 August 4

### ABSTRACT

We show the existence of two families of short gamma-ray bursts (GRBs), both originating from the merger of binary neutron stars (NSs): family-1 with  $E_{\text{iso}} < 10^{52}$  erg, leading to a massive NS as the merged core, and family-2 with  $E_{\text{iso}} > 10^{52}$  erg, leading to a black hole (BH). Following the identification of the prototype GRB 090227B, we present the details of a new example of family-2 short burst: GRB 140619B. From the spectral analysis of the early  $\sim 0.2$  s, we infer an observed temperature  $kT = (324 \pm 33)$  keV of the  $e^+e^-$ -plasma at transparency (P-GRB), a theoretically derived redshift  $z = 2.67 \pm 0.37$ , a total burst energy  $E_{e^+e^-}^{\text{tot}} = (6.03 \pm 0.79) \times 10^{52}$  erg, a rest-frame peak energy  $E_{p,i} = 4.7$  MeV, and a baryon load  $B = (5.52 \pm 0.73) \times 10^{-5}$ . We also estimate the corresponding emission of gravitational waves. Two additional examples of family-2 short bursts are identified: GRB 081024B and GRB 090510, remarkable for its well determined cosmological distance. We show that marked differences exist in the nature of the afterglows of these two families of short bursts: family-2 bursts, leading to BH formation, consistently exhibit high energy emission following the proper-GRB emission; family-1 bursts, leading to the formation of a massive NS, should never exhibit high energy emission. We also show that both the families fulfill an  $E_{p,i}-E_{\text{iso}}$  relation with slope  $\gamma = 0.59 \pm 0.07$  and a normalization constant incompatible with the one for long GRBs. The observed rate of such family-2 events is  $\rho_0 = (2.1_{-1.4}^{+2.8}) \times 10^{-4}$  Gpc<sup>-3</sup> yr<sup>-1</sup>.

*Key words:* gamma-ray burst: general

### 1. INTRODUCTION

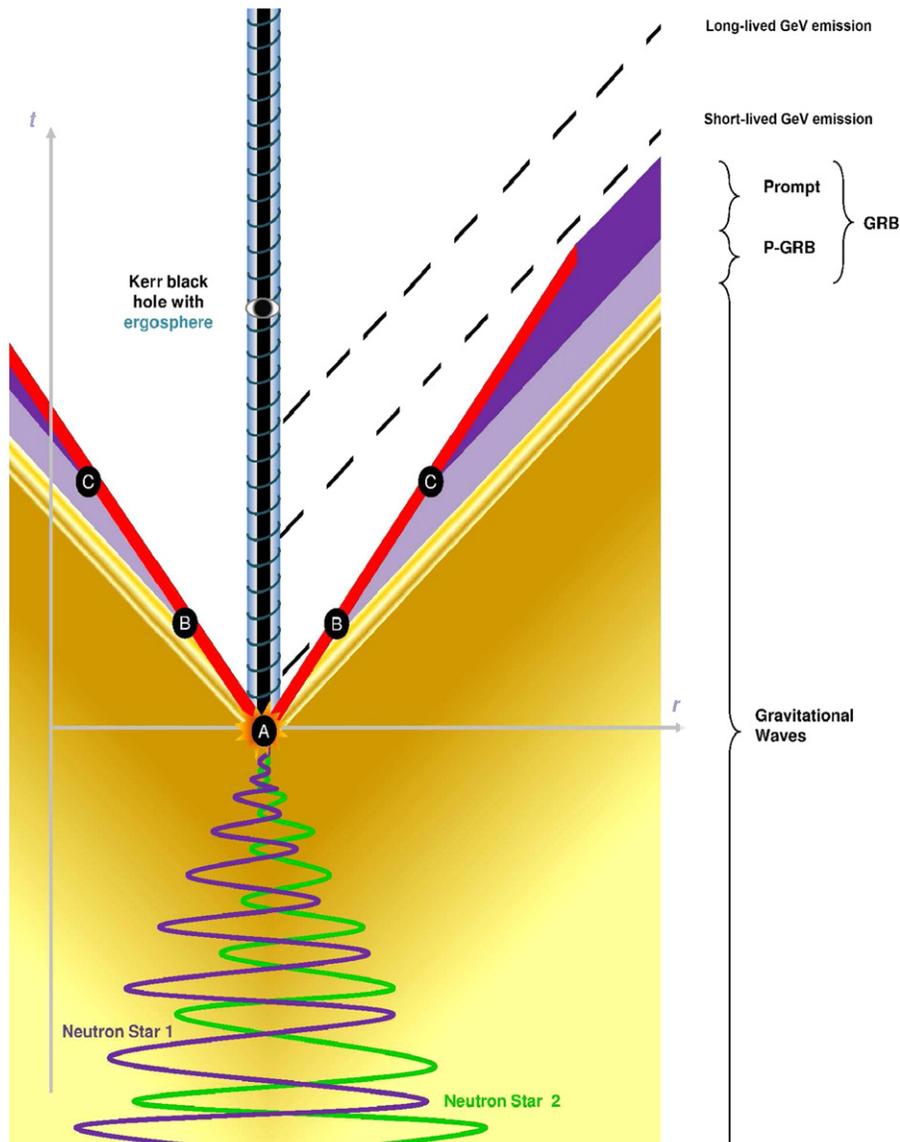
The phenomenological classification of gamma-ray bursts (GRBs) based on their prompt emission observed  $T_{90}$  durations defines “long” and “short” bursts which are, respectively, longer or shorter than  $T_{90} = 2$  s (Dezalay et al. 1992; Klebesadel 1992, pp. 161–168; Kouveliotou et al. 1993; Tavani 1998). Short GRBs have been often indicated as originating from binary neutron star (NS) mergers (see, e.g., Goodman 1986; Paczynski 1986; Eichler et al. 1989; Narayan et al. 1991; Meszaros & Rees 1997; Rosswog et al. 2003; Lee et al. 2004; Berger 2014).

An ample literature exists of short GRBs with a measured redshift, isotropic burst energy  $E_{\text{iso}} < 10^{52}$  erg and rest-frame spectral peak energy  $E_{p,i} < 2$  MeV (see, e.g., Berger 2014 and references therein). Thanks to extensive data provided by the *Swift*-XRT instrument (Burrows et al. 2005), it is possible to observe the long lasting X-ray afterglow of these short bursts to identify their host galaxies and to compute their cosmological redshifts. They have been observed in both early- and late-type galaxies with older stellar population ages (see, e.g., Berger 2014 for details), and at systematically larger radial offsets from their host galaxies than long GRBs (Sahu et al. 1997; van Paradijs et al. 1997; Bloom et al. 2006; Troja et al. 2008; Fong et al. 2010; Berger 2011; Kopač et al. 2012). None of these afterglows appears to have the specific power law signature in the X-ray luminosity when computed in the source rest-frame, as found in some long GRBs (see, e.g., Ruffini et al. 2014).

In the meantime, considerable progress has been obtained in the theoretical understanding of the equilibrium configuration of NSs, in their mass–radius relation (see Figure 2 in Section 2), and especially in the theoretical determination of the value of the NS critical mass for gravitational collapse  $M_{\text{crit}}^{\text{NS}}$  (Rotondo et al. 2011; Rueda et al. 2011; Belvedere et al. 2012). This has led to a theoretical value  $M_{\text{crit}}^{\text{NS}} = 2.67 M_{\odot}$  (Belvedere et al. 2012). Particularly relevant to this determination has been the conceptual change of paradigm of imposing global charge neutrality (Belvedere et al. 2012) instead of the traditional local charge neutrality (LCN) still applied in the current literature (see, e.g., Haensel et al. 2007 and references therein).

Similarly, noteworthy progress has been achieved in the determination of the masses of galactic binary pulsars. Of the greatest relevance has been the direct observation of NS masses larger than  $2 M_{\odot}$  (see Antoniadis et al. 2013 and Section 2). In the majority of the observed cases of binary NSs the sum of the NS masses,  $M_1 + M_2$ , is indeed smaller than  $M_{\text{crit}}^{\text{NS}}$  and, given the above determination of the NS critical mass, their coalescence will never lead to a black hole (BH) formation (see Figure 3 in Section 2). This of course offers a clear challenge to the traditional assumption that all short GRBs originate from BH formation (see, e.g., Berger 2014 and references therein).

Motivated by the above considerations, we propose in this article the existence of two families of short GRBs, both originating from NS mergers: the difference between these two families depends on whether the total mass of the merged core



**Figure 1.** Space–time diagram of family-2 short GRBs. The orbital separation between the two NSs decreases due to the emission of GWs, until the merging occurs and a family-2 short GRB is emitted. Following the fireshell model (see Section 3): (A) vacuum polarization occurs while the event horizon is formed and a fireshell of  $e^+e^-$  plasma self-accelerates radially outwards; (B) the fireshell, after engulfing the baryons, keeps self-accelerating and reaches the transparency when the P-GRB is emitted; (C) the accelerated baryons interact with the local CBM giving rise to the prompt emission. The remnant of the merger is a Kerr BH. The accretion of a small (large) amount of orbiting matter onto the BH can lead to the short lived but very energetic 0.1–100 GeV emission observed in GRB 081024B, GRB 090510, and GRB 140619B. The absence of such an emission in GRB 090227B is due to the absence of observations of *Fermi*-LAT.

is smaller or larger than  $M_{\text{crit}}^{\text{NS}}$ . We assume that family-1 coincides with the above mentioned less energetic short GRBs with  $E_{\text{iso}} < 10^{52}$  erg and the coalescence of the merging NSs leads to a massive NS as the merged core. We assume that family-2 short bursts with  $E_{\text{iso}} > 10^{52}$  erg originate from a merger process leading to a BH as the merged core. The presence of the BH allows us to address the GRB nature within the fireshell model (Ruffini et al. 2001a, 2001b, 2001c) leading to specific signatures in the luminosity, spectra and time variability observed in two very different components: the proper-GRB (P-GRB) and the prompt emission (see Section 3). The prototype is GRB 090227B, which we already analyzed within the fireshell model in Muccino et al. (2013). We also assume that the BH gives rise to the short-lived ( $\lesssim 10^2$  s in the observer frame) and very energetic GeV emission which has been found to be present in all these family-2 short GRBs,

when *Fermi*-LAT data are available. This article is mainly dedicated to giving the theoretical predictions and the observational diagnostics to support the above picture.

In Section 4 we recall the results obtained in the case of the prototype of family-2 short GRBs: GRB 090227B (Muccino et al. 2013). The analysis of its P-GRB emission led to a particularly low value of the baryon load,  $B \sim 10^{-5}$ , as well as to the prediction of the distance corresponding to a redshift  $z = 1.61$ , and consequently to  $E_{e^+e^-}^{\text{tot}} = 2.83 \times 10^{53}$  erg. From the analysis of the spectrum and the light curve of the prompt emission we inferred an average circumburst medium (CBM) density  $\langle n_{\text{CBM}} \rangle \sim 10^{-5} \text{ cm}^{-3}$  typical of galactic halos of GRB host galaxies.

In Section 5 we summarize the observations of a second example of such family-2 short bursts, GRB 140619B, and our data analysis from 8 keV up to 100 GeV. We also point out the

lack of any observed X-ray afterglow following the prompt emission (Maselli & D’Avanzo 2014).

In Section 6 we address GRB 140619B within the fireshell model and compare and contrast the results with those of the prototype, GRB 090227B (Muccino et al. 2013). In Section 6.1, from the fireshell equations of motion, we theoretically estimate and predict the value of the redshift of the source,  $z = 2.67 \pm 0.37$ . Consequently, we derive the burst energy  $E_{\text{iso}} > 10^{52}$  erg and the value of the baryon load  $B \sim 10^{-5}$ . In Section 6.2 we infer an average density of the CBM ( $n_{\text{CBM}} \sim 10^{-5} \text{ cm}^{-3}$  from fitting the prompt emission light curve and spectra. This parameter is typical of the galactic halo environment and further confirms a NS–NS merger as the progenitor for GRB 140619B (see Section 6.3 and Figure 1).

In Section 7 we discuss the possibility for Advanced LIGO to detect the emission of gravitational waves (GWs) from such a binary NS progenitor. From the dynamics of the above system, the total energy emitted in GW radiation corresponds to  $E_{\text{GW}}^T = 7.42 \times 10^{52}$  erg, computed during the entire inspiral phase all the way up to the merger. This gives a signal below the sensitivity of the Advanced LIGO interferometer.

In Section 8 we focus on the short-lived ( $\Delta t \approx 4$  s) but significant 0.1–100 GeV emission (see Figure 1). We first address the issue of whether this is a peculiarity of GRB 140619B, or whether the GeV emission can be considered to be a common feature of all these family-2 short GRBs. We first return to GRB 090227B to see how to explain the absence of observations of the GeV emission from this source, and we find a simple reason: GRB 090227B was outside the nominal LAT field of view (FOV, see Ackermann et al. 2013, and Section 4). We then turn our attention to another source, GRB 090510, which presents many of the common features of the family-2 short GRBs. Especially noteworthy is the presence of a high energy GeV emission lasting  $\sim 10^2$  s, much longer than the one of GRB 140619B. The presence of an X-ray afterglow in GRB 090510 is fortunate and particularly important, though lacking a scaling law behavior (Ruffini et al. 2014), since it has allowed the optical identification of the source and the determination of its distance and its cosmological redshift  $z = 0.903$ . The corresponding isotropic energy and intrinsic peak spectral energy are, respectively,  $E_{\text{iso}} > 10^{52}$  erg and  $E_{p,i} = (7.89 \pm 0.76)$  MeV, typical again of family-2 short bursts. We then compare and contrast this high energy emission and their corresponding X-ray emissions in the family-2 short GRB 140619B and GRB 090510 with the afterglow of the family-1 short GRBs (see Figure 13 and Berger 2014).

In Section 9 we give an estimate for the rate of the family-2 short GRBs.

In Section 10 we discuss the existence of the new  $E_{p,i}$ – $E_{\text{iso}}$  relation for all short GRBs introduced by Zhang et al. (2012) and Calderone et al. (2015), with a power-law similar to the one of the Amati relation (Amati et al. 2008) for long GRBs, but with a different amplitude. Finally we draw our conclusions.

## 2. MOTIVATION FROM GALACTIC BINARY NS AND NS THEORY

Recent theoretical progress has been achieved in the understanding of the NS equation of state and equilibrium configuration and of the value of its critical mass  $M_{\text{crit}}^{\text{NS}}$ . In Rotondo et al. (2011) it has been shown to be impossible to

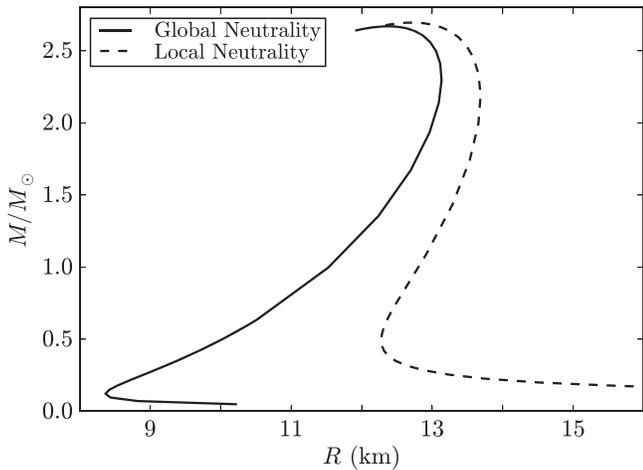
impose the LCN condition on a self-gravitating system of degenerate neutrons, protons, and electrons in  $\beta$ -equilibrium within the framework of relativistic quantum statistics and the Einstein–Maxwell equations. The equations of equilibrium of NSs, taking into account strong, weak, electromagnetic, and gravitational interactions in general relativity and the equilibrium conditions based on the Einstein–Maxwell–Thomas–Fermi equations along with the constancy of the general relativistic Fermi energies of particles, the “Klein potentials,” throughout the configuration have been presented in Rueda et al. (2011) and Belvedere et al. (2012), where a theoretical estimate of  $M_{\text{crit}}^{\text{NS}} \approx 2.67 M_{\odot}$  has been obtained. The implementations of the above results by considering the equilibrium configurations of slowly rotating NSs by using the Hartle formalism has been presented in Belvedere et al. (2014a). Then in Rueda et al. (2014) a detailed study was made of the transition layer between the core and crust of NSs at the nuclear saturation density, and its surface tension and Coulomb energy have been calculated. A comprehensive summary of these results for both static and uniformly rotating NSs is discussed in Belvedere et al. (2014b). The absolute upper limit on the angular momentum of a rotating NS fulfilling the above microscopical conditions has been obtained in Cipolletta et al. (2015).

A vast number of tests have been performed in fitting the data of pulsars (Deneva et al. 2012; Lattimer 2012; Antoniadis et al. 2013; Kramer 2014). In particular, the high value of the recently measured mass of PSR J0348+0432,  $M = (2.01 \pm 0.04) M_{\odot}$  (Antoniadis et al. 2013), favors stiff nuclear equations of state, like the one adopted in Belvedere et al. (2012) based on relativistic nuclear mean field theory á la Boguta & Bodmer (1977), which leads to the above theoretical estimate of  $M_{\text{crit}}^{\text{NS}}$  (see also Figure 2). This value is supported by the above observational constraints, and in any case, is well below the absolute upper limit of  $3.2 M_{\odot}$  for a non-rotating NS (Rhoades & Ruffini 1974).

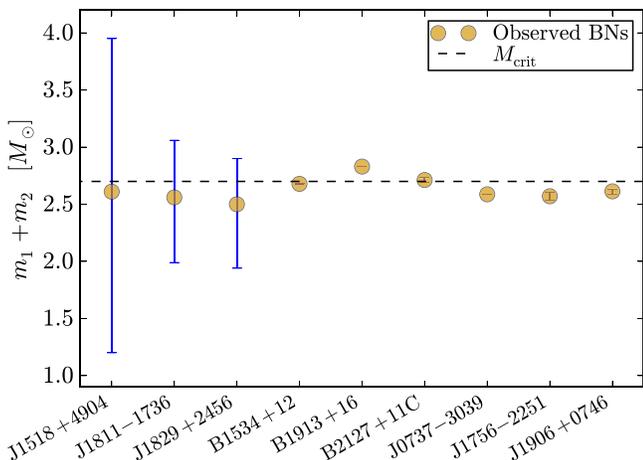
If we turn to the binary NSs within our Galaxy (see Figure 3) we notice that only in a subset of them is the total mass of the components larger than  $M_{\text{crit}}^{\text{NS}}$  and can lead to a BH in their merging process.<sup>6</sup>

Given this general understanding, we have identified the characteristic properties of family-2 short bursts, whose prototype was identified in GRB 090227B (Muccino et al. 2013). Equally important has been the identification of the observed characteristic features of family-1 short GRBs which will be discussed in the following sections.

<sup>6</sup> During the refereeing process, an approach by Fryer et al. (2015) based on a combination of binary NS merger nuclear physics models and population synthesis appeared. They infer that for a maximum nonrotating NS mass of  $M_{\text{crit}}^{\text{NS}}$  above 2.3–2.4  $M_{\odot}$ , less than 4% of the NS mergers produces short GRBs by gravitational collapse to a BH. Here we go one step further by indicating the theoretical predictions characterizing short GRBs originating from the massive NS formation (family-1) and the ones originating from BH formation (family-2). We indicate: (a) the specific spectral features, (b) the presence of the GeV emission originating from the BH, and (c) the fulfillment of the  $E_{p,i}$ – $E_{\text{iso}}$  relation (see Zhang et al. 2012; Calderone et al. 2015, and Section 10). The paper by Fryer et al. (2015) was followed by Lawrence et al. (2015) where the authors examine the value of  $M_{\text{crit}}^{\text{NS}}$  for a family of equations of state and concluded that a reasonable fraction of double NS mergers may produce neither short GRBs nor BHs. Here we again go one step further by indicating that in the case of a merged core with a mass smaller than  $M_{\text{crit}}^{\text{NS}}$  leading to a massive NS, a less energetic short GRB with a softer emission tail indeed occurs (family-1 short bursts). We show also that these short GRBs fulfill the above  $E_{p,i}$ – $E_{\text{iso}}$  relation (see Section 10).



**Figure 2.** Mass–radius relation obtained with the local and the new global neutrality equilibrium configurations, by applying the NL3 nuclear model. Figure reproduced from Belvedere et al. (2012).



**Figure 3.** Plot of the binary NSs with known total masses ( $M_1 + M_2$ , in solar masses) and the corresponding uncertainties. The horizontal dashed line marks the critical NS mass of  $2.67 M_\odot$  (Belvedere et al. 2012). Systems beyond this value lead to BH formation. Masses taken from Zhang et al. (2011) and Antoniadis (2014).

The crucial role of  $M_{\text{crit}}^{\text{NS}}$  has been also shown in the corresponding analysis of long GRBs in distinguishing between the two different families (Ruffini et al. 2015) in the induced gravitational collapse paradigm (Izzo et al. 2012a; Rueda & Ruffini 2012; Fryer et al. 2014).

### 3. THE FIRESHELL MODEL

It is well known that the majority of the astrophysical community working on GRBs envisages the spectral and temporal analysis of both short and long GRBs considering their whole emission as a single event (see, e.g., Ackermann et al. 2013). This picture follows the conceptual framework of the “fireball model” (see, e.g., Sari et al. 1998; Piran 2005; Meszaros 2006, and reference therein).

The “fireshell model” (Ruffini et al. 2001a, 2001b, 2001c) has instead addressed a specific time-resolved spectral analysis leading to distinct signatures and to the identification of different astrophysical regimes within the same GRB (see, e.g., Izzo et al. 2010; Izzo et al. 2012b; Muccino et al. 2013; Ruffini et al. 2013 and references therein). This has led to introduction

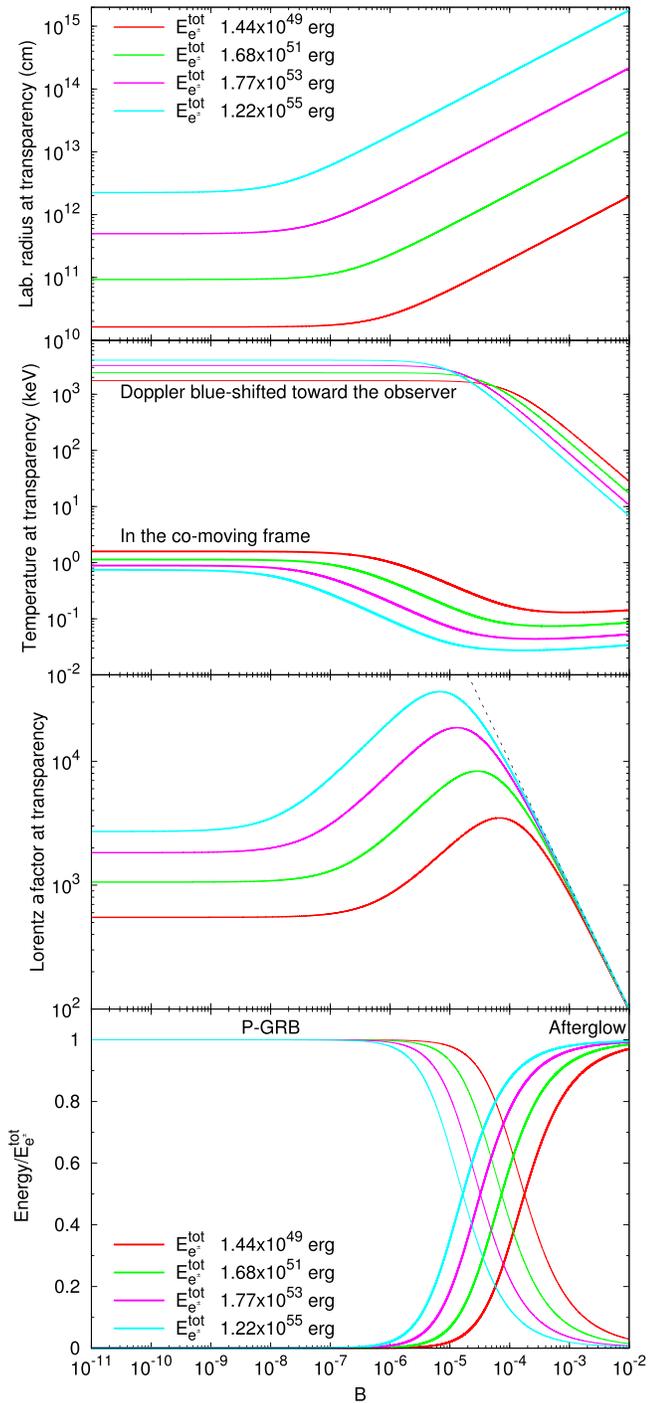
of the concept of binary mergers of NS–NS and of FeCO–NS together with a set of new paradigms in order to describe the complexity of GRB phenomena within a “Cosmic-Matrix” approach (Ruffini 2015a).

In the fireshell model (Ruffini et al. 2001a, 2001b, 2001c) GRBs originate from an optically thick  $e^+e^-$  plasma (Damour & Ruffini 1975; Ruffini & Xue 2008; Ruffini et al. 2010) during the gravitational collapse to a BH. Such an  $e^+e^-$  plasma is confined to an expanding shell and reaches thermal equilibrium almost instantaneously (Aksenov et al. 2007). The annihilation of these pairs occurs gradually, while the expanding shell, called the *fireshell*, self-accelerates up to ultra relativistic velocities (Ruffini et al. 1999) and engulfs the baryonic matter (of mass  $M_B$ ) left over in the process of collapse. The baryon load thermalizes with the pairs due to the large optical depth (Ruffini et al. 2000).

Assuming spherical symmetry of the system, the dynamics in the optically thick phase is fully described by only two free initial parameters: the total energy of the plasma  $E_{e^+e^-}^{\text{tot}}$  and the baryon load  $B$  (Ruffini et al. 2000). Only solutions with  $B \leq 10^{-2}$  are characterized by regular relativistic expansion; for  $B \geq 10^{-2}$  turbulence and instabilities occur (Ruffini et al. 2000). The fireshell continues to self-accelerate until it reaches the transparency condition and a first flash of thermal radiation, the P-GRB, is emitted (Ruffini et al. 2001b). The radius  $r_{\text{tr}}$  at which the transparency occurs, the theoretical temperature (blueshifted toward the observer  $kT_{\text{blue}}$ ), the Lorentz factor  $\Gamma_{\text{tr}}$ , as well as the amount of the energy emitted in the P-GRB are functions of  $E_{e^+e^-}^{\text{tot}}$  and  $B$  (see, e.g., Ruffini et al. 2001b; Ruffini et al. 2009, and Figure 4).

After transparency, the residual expanding plasma of leptons and baryons collides with the CBM giving rise to multi-wavelength emission: the prompt emission. Assuming the fully radiative condition, the structures observed in the prompt emission of a GRB are described by two quantities associated with the environment: the CBM density profile  $n_{\text{CBM}}$ , which determines the temporal behavior of the light curve, and the fireshell surface filling factor  $\mathcal{R} = A_{\text{eff}}/A_{\text{vis}}$ , in which  $A_{\text{eff}}$  is the effective emitting area of the fireshell, and  $A_{\text{vis}}$  is its total visible area (Ruffini et al. 2002, 2005). This second parameter takes into account the inhomogeneities in the CBM and its filamentary structure (Ruffini et al. 2004).

The emission process of the collision between the baryons and the CBM is described in the comoving frame of the shell as a modified blackbody (BB) spectrum. This spectrum is obtained by the introduction of an additional phenomenological parameter  $\alpha$  which characterizes the departure of the slope of the low energy part of the comoving spectrum from the purely thermal one (see Patricelli et al. 2012, for details). The nonthermal spectral shape of the observed GRB is then produced by the convolution of a very large number of modified thermal spectra with different temperatures and different Lorentz and Doppler factors. This convolution is performed over the surfaces of constant arrival time for photons at the detector (EQuiTemporal Surfaces, EQTS, Bianco & Ruffini 2005a, 2005b), encompassing the total observation time. The observed hard-to-soft spectral variation comes out naturally from the decrease with time of the comoving temperature and of the bulk Lorentz  $\Gamma$  factor. This effect is amplified by the curvature effect due to the EQTS which produces the observed time lag in the majority of the GRBs.



**Figure 4.** Main quantities of the fireshell model at transparency for selected values of  $E_{e^+e^-}^{\text{tot}}$ : the radius in the laboratory frame, the temperatures of the plasma in the co-moving frame and blueshifted toward the observer, the Lorentz  $\Gamma$  factor, and the fraction of energy radiated in the P-GRB and in the prompt emission as functions of  $B$ .

The canonical GRB light curve within the fireshell model is then characterized by a first (mainly thermal) emission due to the transparency of the  $e^+e^-$ -photon-baryon plasma, the P-GRB. A multi-wavelength emission, the prompt emission, follows due to the collisions between the accelerated baryons and the CBM.

The fireshell model has originally described the process of vacuum polarization due to the overcritical electromagnetic

field occurring at the moment of BH formation (Damour & Ruffini 1975). The formalism has been developed by considering a large number of relativistic quantum effects in the electrodynamics proposed for the NS crust (Belvedere et al. 2012, 2014a; Rueda et al. 2014), as well as on quantum-electrodynamics processes ongoing in the gravitational collapse (Han et al. 2012; Ruffini & Xue 2013). This has led to the results summarized in Figure 4.

The first description of the  $e^+e^-$  plasma within the fireshell model was performed under the simplified assumption of spherical symmetry (the dyadosphere; see, e.g., Preparata et al. 1998). The corresponding structure in the axially symmetric Kerr-Newman geometry has been considered (the dyadotorus; see, e.g., Cherubini et al. 2009; Ruffini 2009) and could possibly be tested.

The general formalism of the fireshell model can also be applied to any optically thick  $e^+e^-$  plasma in the presence of a baryon load, like the one created during the merging of binary NSs from  $\nu\bar{\nu} \rightarrow e^+e^-$  (see, e.g., Salmonson & Wilson 2002 and references therein).

The P-GRB addresses the fully relativistic fundamental physics aspects of the model, in particular the acceleration process of the  $e^+e^-$ -baryon plasma, the collapsing NS quantum-electrodynamics, and the BH physics. The prompt emission addresses the conceptually simpler problem of the interaction of the accelerated baryons with the CBM, which does not allow nor require, by its own nature, a detailed description.

#### 4. SUMMARY OF THE RESULTS FOR GRB 090227B: THE PROTOTYPE OF THE FAMILY-2 SHORT GRBS

GRB 090227B is a bright short burst with an overall emission lasting  $\sim 0.9$  s and total fluence of  $3.79 \times 10^{-5}$  erg  $\text{cm}^{-2}$  in the energy range 8 keV–40 MeV. This burst was significantly detected only in the LAT Low Energy (LLE) data since it was outside the nominal LAT FOV (Ackermann et al. 2013). However, only one transient-class event with energy above 100 MeV has been associated with the GRB (Ackermann et al. 2013).

The time-resolved spectral analysis on the time scale as short as 16 ms, made possible by the *Fermi*-GBM (Meegan et al. 2009), has allowed the identification of the P-GRB in the early 96 ms of emission. The corresponding thermal component has a temperature  $kT = (517 \pm 28)$  keV (see the upper plots of Figure 9 in Muccino et al. 2013). The subsequent emission, fit by a Band function (see lower plots of Figure 9 in Muccino et al. 2013), has been identified with the prompt emission.

Due to the absence of an optical identification, a direct measurement of the cosmological redshift was not possible. From the temperature and flux of the P-GRB thermal component it was possible to derive (see Figure 4) a theoretical cosmological redshift  $z = 1.61 \pm 0.14$ , as well as the baryon load  $B = (4.13 \pm 0.05) \times 10^{-5}$ , the total plasma energy  $E_{e^+e^-}^{\text{tot}} = (2.83 \pm 0.15) \times 10^{53}$  erg, and the extremely high Lorentz  $\Gamma$  factor at transparency  $\Gamma_{\text{tr}} = (1.44 \pm 0.01) \times 10^4$  (see Section 4.1 in Muccino et al. 2013). Consequently, an average CBM number density  $\langle n_{\text{CBM}} \rangle = (1.90 \pm 0.20) \times 10^{-5}$   $\text{cm}^{-3}$  has been determined which is typical of galactic halos where NS–NS mergers migrate, owing to natal kicks imparted to the binaries at birth (see, e.g., Narayan et al. 1992; Bloom et al. 1999; Fryer et al. 1999; Belczynski et al. 2006; Berger 2014).

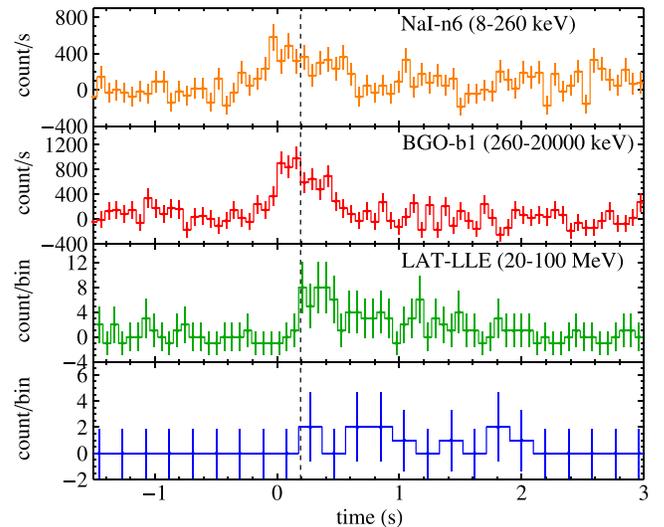
In Muccino et al. (2013) it was concluded that the progenitor of GRB 090227B is a binary NS. For simplicity and as a lower limit, the masses of the two NS have been assumed to be the same, e.g.,  $M_1 = M_2 = 1.34 M_\odot$ , so that the total merged core mass is  $>M_{\text{crit}}^{\text{NS}}$  and therefore a BH is formed. This conclusion was drawn in view of the large total energy,  $E_{e^+e^-}^{\text{tot}} = 2.83 \times 10^{53}$  erg. Correspondingly, the energy emitted via GWs,  $\sim 9.7 \times 10^{52}$  erg, has been estimated in Oliveira et al. (2014).

## 5. OBSERVATIONS AND DATA ANALYSIS OF GRB 140619B

At 11:24:40.52 UT on 2014 June 19, the *Fermi*-GBM detector (Connaughton et al. 2014) triggered and located the short and hard burst GRB 140619B (trigger 424869883/140619475). The on-ground calculated location, using the GBM trigger data, was R.A.(J2000) = 08<sup>h</sup>54<sup>m</sup> and decl. (J2000) =  $-3^\circ 42'$ , with an uncertainty of  $5^\circ$  (statistical only). The location of this burst was  $32^\circ$  from the LAT boresight at the time of the trigger, and the data from the *Fermi*-LAT showed a significant increase in the event rate (Kocevski et al. 2014). The burst was also detected by *Suzaku*-WAM (Iwakiri et al. 2014), showing a single pulse with a duration of  $\sim 0.7$  s (50 keV–5 MeV). The analysis from 48.7 to 71.6 ks after the GBM trigger by the *Swift*-XRT instrument in the FOV of the *Fermi*-GBM and LAT, was completely in Photon Counting mode (Maselli & D’Avanzo 2014). No bright X-ray afterglow was detected within the LAT error circle. This set an upper limit on the energy flux in the observed 0.3–10 keV energy band of  $\approx 9.24 \times 10^{-14}$  erg/(cm<sup>2</sup> s), assuming a photon index  $\gamma = 2.2$ . Therefore, no optical follow-up was possible and thus the redshift of the source is unknown.

We have analyzed the *Fermi*-GBM and LAT data in the energy range 8 keV–40 MeV and 20 MeV–100 GeV, respectively. We have downloaded the GBM TTE (Time-Tagged Events) files,<sup>7</sup> suitable for short or highly structured events, and analyzed them by using the RMFIT package.<sup>8</sup> The LLE data<sup>9</sup>, between 20–100 MeV, and the high energy data<sup>10</sup>, between 100 MeV–100 GeV, were analyzed by using the Fermi-science tools.<sup>11</sup> In Figure 5 we have reproduced the 64 ms binned GBM light curves corresponding to detectors NaI-n6 (8–260 keV, top panel) and BGO-b1 (260 keV–20 MeV, second panel), the 64 ms binned LLE light curve (20–100 MeV, third panel) and the 192 ms binned high-energy channel light curve (0.1–100 GeV, bottom panel). All the light curves are background subtracted. The NaI-n6 light curve shows a very weak signal, almost at the background level, while the BGO-b1 signal is represented by a short hard pulse, possibly composed by two sub-structures, with a total duration of  $T_{90} \approx 0.7$  s. The vertical dashed line in Figure 5 represents the on-set of both LAT light curves, i.e.,  $\sim 0.2$  s after the GBM trigger. In principle, this allows us to determine the time interval within which the P-GRB emission takes place.

We have subsequently performed the time-integrated and time-resolved spectral analyses focused on the GBM data in the energy range 8 keV–40 MeV.



**Figure 5.** Background subtracted light curves of GRB 140619B from various detectors in various energy bands. From the top to the bottom panel: the 64 ms binned light curves from the NaI-n6 (8–260 keV, top panel) and BGO-b1 (260 keV–20 MeV, second panel) detectors, the 64 ms binned LLE light curve (20–100 MeV, third panel), and the 192 ms binned high-energy channel light curve (100 MeV–100 GeV, bottom panel).

### 5.1. Time-integrated Spectral Analysis

We have performed a time-integrated spectral analysis in the time interval from  $T_0 - 0.064$  s to  $T_0 + 0.640$  s, which corresponds to the  $T_{90}$  duration of the burst. We have indicated the trigger time by  $T_0$  and have considered the following spectral models: Comptonization (Compt) and a Band function (Band et al. 1993). The corresponding plots are shown in Figure 6 and the results of the fits are listed in Table 1. From a statistical point of view, the Compt model provides the best fit to the data. In fact the Band function, which has an additional parameter with respect to the Compt model, improves the fit by only  $\Delta\text{C-STAT} = 2.53$ , where  $\Delta\text{C-STAT}$  is the difference between the two C-STAT values of the Compt and Band models. If we consider  $\Delta\text{C-STAT}$  as a  $\chi^2$  variable for the change in the number of the model parameters  $\Delta n$  (in this case  $\Delta n = 1$ ), and assuming that the Compt model is nested within the Band model,<sup>12</sup> we conclude that the Band model improves the fit only at the 89% significance level, and anyway less than  $2\sigma$ . Therefore it is not enough to reject the Compt model. The most interesting feature of the Compt model consists of its low-energy index, which is consistent with  $\alpha \sim 0$ . We proceed now to a time-resolved analysis to investigate the possibility that in the early phases of the prompt emission the spectrum is consistent with a BB spectrum, i.e.,  $\alpha \approx 1$ , which corresponds to the signature of P-GRB emission.

### 5.2. Time-resolved Spectral Analysis

We performed the time-resolved spectral analysis by selecting time intervals with fluences larger than  $\approx 10^{-6}$  erg cm<sup>-2</sup> in order to collect enough photons. Consequently, we have selected two time intervals that correspond to the main spike and the less intense structure (see the BGO-b1 light curve in Figure 6). The first time interval, from  $T_0$  to

<sup>7</sup> <ftp://legacy.gsfc.nasa.gov/fermi/data/gbm/bursts>

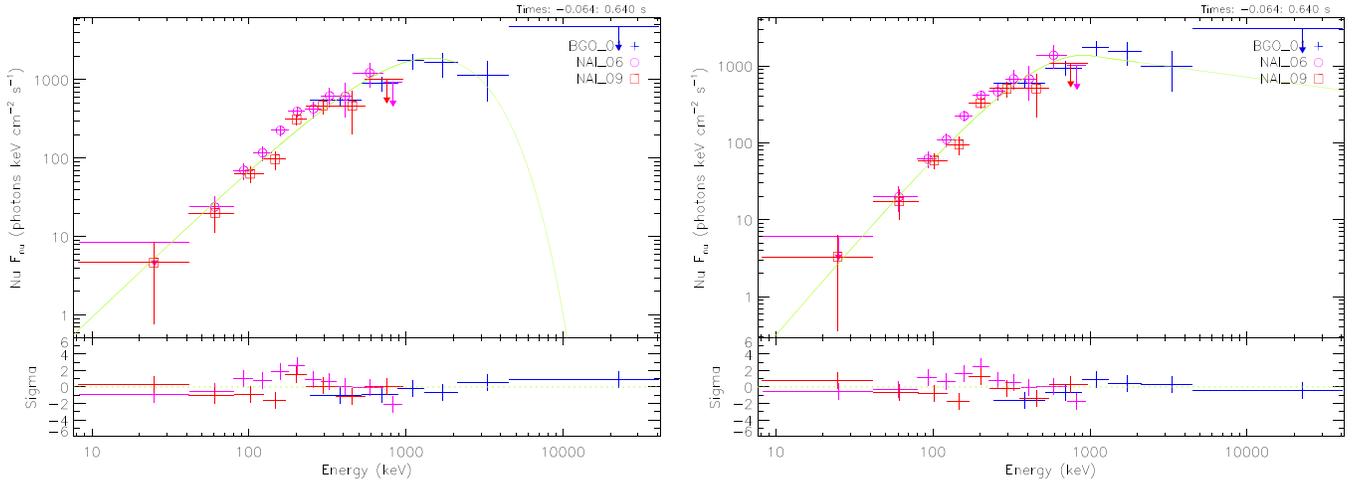
<sup>8</sup> [http://Fermi.gsfc.nasa.gov/ssc/data/analysis/rmfit/vc\\_rmfit\\_tutorial.pdf](http://Fermi.gsfc.nasa.gov/ssc/data/analysis/rmfit/vc_rmfit_tutorial.pdf)

<sup>9</sup> [http://fermi.gsfc.nasa.gov/ssc/observations/types/grbs/lat\\_grbs/](http://fermi.gsfc.nasa.gov/ssc/observations/types/grbs/lat_grbs/)

<sup>10</sup> <http://fermi.gsfc.nasa.gov/cgi-bin/ssc/LAT/LATDataQuery.cgi>

<sup>11</sup> <http://fermi.gsfc.nasa.gov/ssc/data/analysis/documentation/Cicerone/>

<sup>12</sup> The Compt model can be considered a particular case of the Band model with  $\beta \rightarrow -\infty$ .



**Figure 6.** Combined NaI-n6, n9+BGO-b1  $\nu F_\nu$  spectra of GRB 140619B in the  $T_{90}$  time interval. The fit using the Compt spectral model is shown on the left, while the Band model fit is on the right.

**Table 1**  
Summary of the Time-integrated ( $T_{90}$ ) and Time-resolved ( $\Delta T_1$  and  $\Delta T_2$ ) Spectral Analyses

$\Delta T$	Model	$K$ (ph keV $^{-1}$ cm $^{-2}$ s $^{-1}$ )	$kT$ (keV)	$E_p$ (keV)	$\alpha$	$\beta$	$F_{\text{tot}}$ (erg cm $^{-2}$ s $^{-1}$ )	C-STAT/DOF
$T_{90}$	Compt	$(7.7 \pm 1.1) \times 10^{-3}$	...	$1456 \pm 216$	$-0.09 \pm 0.18$	...	$(5.75 \pm 0.75) \times 10^{-6}$	365.09/346
	Band	$(7.8 \pm 1.3) \times 10^{-3}$	...	$908 \pm 199$	$-0.38 \pm 0.37$	$-2.28 \pm 0.31$	$(7.4 \pm 1.8) \times 10^{-6}$	362.56/345
$\Delta T_1$	Compt	$(6.3 \pm 2.0) \times 10^{-3}$	...	$1601 \pm 287$	$0.26 \pm 0.32$	...	$(9.4 \pm 1.6) \times 10^{-6}$	318.92/346
	BB	$(7.5 \pm 2.2) \times 10^{-8}$	$324 \pm 33$	...	...	...	$(8.5 \pm 1.2) \times 10^{-6}$	323.86/347
$\Delta T_2$	Compt	$(7.2 \pm 1.4) \times 10^{-3}$	...	$1283 \pm 297$	$-0.11 \pm 0.26$	...	$(4.38 \pm 0.89) \times 10^{-6}$	391.65/346
	BB	$(3.8 \pm 1.1) \times 10^{-7}$	$156 \pm 15$	...	...	...	$(2.33 \pm 0.28) \times 10^{-6}$	392.23/347

**Note.** In each column are listed, respectively, the time interval  $\Delta T$ , the adopted spectral model, the normalization constant  $K$  of the fitting function, the BB temperature  $kT$ , the peak energy  $E_p$ , the low-energy  $\alpha$  and high-energy  $\beta$  photon indexes, the total energy flux  $F_{\text{tot}}$  in the range 8 keV–40 MeV, and the value of the C-STAT over the number of degrees of freedom (dof).

$T_0 + 0.192$  s, is hereafter referred to as  $\Delta T_1$ , while the subsequent emission, from  $T_0 + 0.192$  s to  $T_0 + 0.640$  s, is designated by  $\Delta T_2$ .

In the  $\Delta T_1$  time interval, to identify the P-GRB, we have performed a spectral analysis by considering the BB and Compt spectral models. The spectra and the corresponding fits are shown in Figure 7 and the best fit parameters are listed in Table 1. As reported in Table 1, the Compt and the BB models are both viable. However, the value of the low-energy index of the Compt model in the  $\Delta T_1$  time interval,  $\alpha = 0.26 \pm 0.32$ , is consistent within three  $\sigma$  with  $\alpha = 1$ , which is the low energy index of a BB. We conclude that the BB model is an acceptable fit to the data and the best “physical model” of the  $\Delta T_1$  time interval and therefore identify it with the P-GRB emission. The corresponding observed temperature is  $kT = (324 \pm 33)$  keV (see Table 1).

We then performed a spectral analysis on the time interval  $\Delta T_2$  to identify the prompt emission. We have again considered the Compt and BB spectral models (see Figure 8 and Table 1). By looking at Figure 8, it is immediately clear that the BB model does not adequately fit the data at energies larger than 1 MeV. Therefore the Compt model is favored. Its low-energy index,  $\alpha = -0.11 \pm 0.26$ , indicates that the spectral energy distribution in the  $\Delta T_2$  time interval is broader than that of the BB model. The Compt model is consistent with the spectral model adopted in the fireshell model and described in Patricelli et al. (2012) for the prompt emission.

In the next section we interpret the above data within the fireshell theoretical framework.

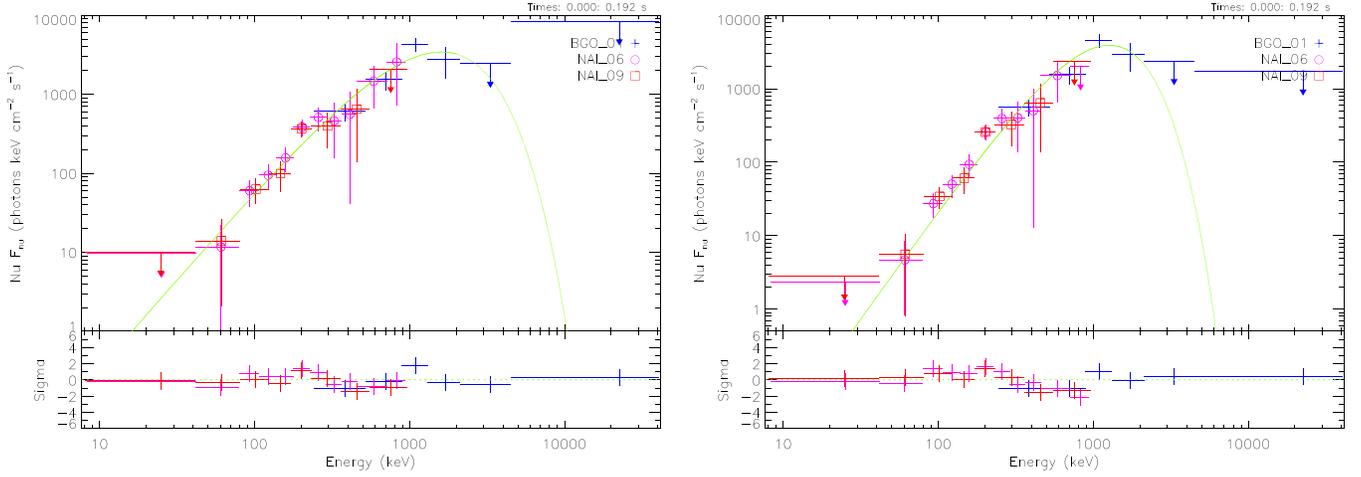
## 6. APPLICATION OF THE FIRESHHELL MODEL TO GRB 140619B

After the P-GRB and the prompt emission identification, we have followed the same analysis described in Muccino et al. (2013) to determine the cosmological redshift, the baryon load and all the other physical quantities characterizing the plasma at the transparency point (see Figure 4). It is appropriate to underline that a remarkable difference between the long and the short GRBs is considered: the P-GRB emission in long GRBs represents on average the 1%–5% of the overall emission (see, e.g., the cases of GRB 970828, Ruffini et al. 2013, and GRB 090618, Izzo et al. 2012b), while in the cases of the short GRBs 090227B and 140619B (see Section 6.1), the P-GRB emissions represent  $\sim 40\%$  of the overall observed fluence.

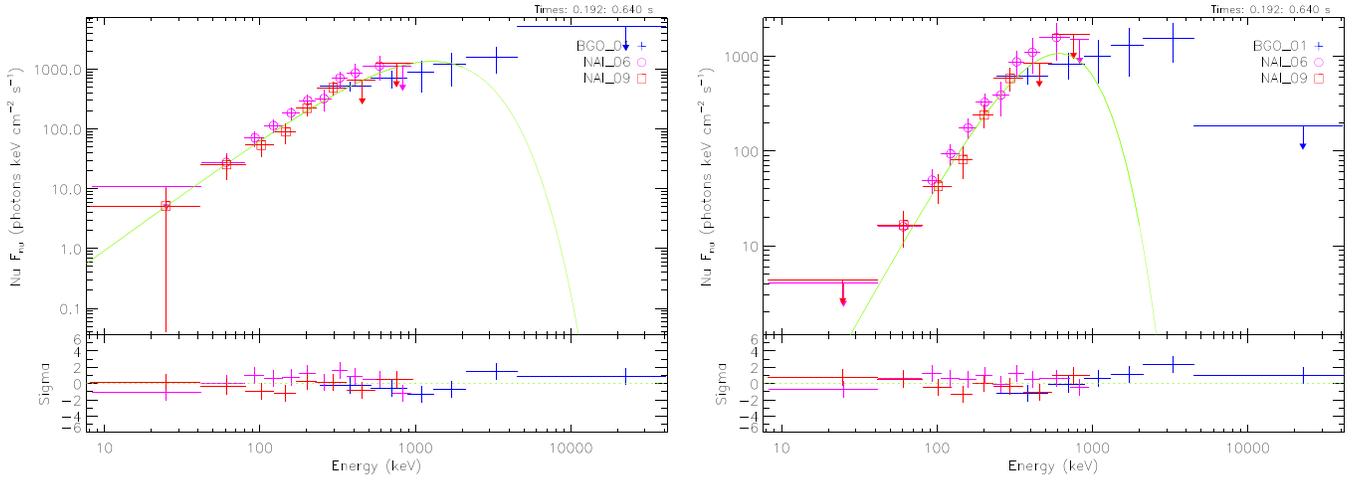
### 6.1. Redshift Estimate in Fireshell Model

From the observed P-GRB and total fluences, respectively,  $S_{\text{BB}} = F_{\text{tot}}(\Delta T_1)\Delta T_1$  and  $S_{\text{tot}} = F_{\text{tot}}(T_{90})T_{90}$  (see values in Table 1), we have estimated the ratio

$$\frac{E_{\text{P-GRB}}}{E_{e^+e^-}^{\text{tot}}} \approx \frac{4\pi d_l^2 S_{\text{BB}}/(1+z)}{4\pi d_l^2 S_{\text{tot}}/(1+z)} = \frac{S_{\text{BB}}}{S_{\text{tot}}} = (40.4 \pm 7.8)\%, \quad (1)$$



**Figure 7.** Same considerations as in Figure 6, in the  $\Delta T_1$  time interval, comparing Compt (left panel) and BB (right panel) models.



**Figure 8.** Same considerations as in Figure 6, in the  $\Delta T_2$  time interval, comparing Compt (left panel) and BB (right panel) models.

where the theoretically computed energy of the P-GRB,  $E_{\text{P-GRB}}$ , has been constrained by the observed thermal emission,  $E_{\text{BB}}$ , and we have imposed  $E_{e^+e^-}^{\text{tot}} \equiv E_{\text{iso}}$ . In Equation (1) the luminosity distance  $d_l$  and the redshift  $z$  of the source do not enter into the final computation.

From the last diagram in Figure 4, it is clear that for the value in Equation (1), we have different possible parameters ( $E_{e^+e^-}^{\text{tot}}$ ,  $B$ ) and for each of them we can determine the corresponding  $kT_{\text{blue}}$  (see the top diagram in Figure 4). Finally, from the ratio between  $kT_{\text{blue}}$  and the observed P-GRB temperature  $kT$ , we can estimate the redshift, i.e.,  $kT_{\text{blue}}/kT = (1+z)$ . To obtain the correct value of  $z$  and then the right parameters [ $E_{e^+e^-}^{\text{tot}}(z)$ ,  $B(z)$ ], we have made use of the isotropic energy formula

$$E_{\text{iso}} = 4\pi d_l^2 \frac{S_{\text{tot}}}{(1+z)} \frac{\int_{1/(1+z)}^{10000/(1+z)} E N(E) dE}{\int_8^{40000} E N(E) dE}, \quad (2)$$

in which  $N(E)$  is the photon spectrum of the burst and the integrals are due to the  $K$ -correction on  $S_{\text{tot}}$  (Schaefer 2007). From the initial constraint  $E_{\text{iso}} \equiv E_{e^+e^-}^{\text{tot}}$ , we have found  $z = 2.67 \pm 0.37$ , which leads to  $B = (5.52 \pm 0.73) \times 10^{-5}$

and  $E_{e^+e^-}^{\text{tot}} = (6.03 \pm 0.79) \times 10^{52}$  erg. All the quantities so determined are summarized in Table 2. The analogy with the prototypical source GRB 090227B, for which we have  $E_{\text{P-GRB}} = (40.67 \pm 0.12)\% E_{e^+e^-}^{\text{tot}}$  and  $B = (4.13 \pm 0.05) \times 10^{-5}$ , is very striking (Muccino et al. 2013).

We now proceed with the analysis of the subsequent emission to derive the properties of the surrounding CBM.

## 6.2. Analysis of the Prompt Emission

Having determined the initial conditions for the fireshell, i.e.,  $E_{e^+e^-}^{\text{tot}} = 6.03 \times 10^{52}$  erg and  $B = 5.52 \times 10^{-5}$ , the dynamics of the system is uniquely established. In particular, we obtain the Lorentz factor at transparency,  $\Gamma_{\text{tr}} = 1.08 \times 10^4$ , and we can simulate the light curve and the spectrum of the prompt emission. To reproduce the pulses observed especially in the BGO-b1 light curve (see Figure 5) we have derived the radial distributions of the CBM number density and of the filling factor  $\mathcal{R}$  around the burst site (see Table 3 and Figure 9). The errors in the CBM number density and in  $\mathcal{R}$  are defined as the maximum possible variation of the parameters to guarantee agreement between the simulated light curve and the observed data. The final simulation of the BGO-b1 light curve (260 keV–40 MeV) is shown in Figure 10.

**Table 2**  
The Results of the Simulation of GRB 090227B in the Fireshell Model

Fireshell Parameter	Value
$E_{e^+e^-}^{\text{tot}}$ (erg)	$(6.03 \pm 0.79) \times 10^{52}$
$B$	$(5.52 \pm 0.73) \times 10^{-5}$
$\Gamma_{\text{tr}}$	$(1.08 \pm 0.08) \times 10^4$
$r_{\text{tr}}$ (cm)	$(9.36 \pm 0.42) \times 10^{12}$
$kT_{\text{blue}}$ (keV)	$(1.08 \pm 0.08) \times 10^3$
$z$	$2.67 \pm 0.37$
$\langle n_{\text{CBM}} \rangle$ ( $\text{cm}^{-3}$ )	$(4.7 \pm 1.2) \times 10^{-5}$

Interestingly, the average CBM number density in GRB 140619B,  $\langle n_{\text{CBM}} \rangle = (4.7 \pm 1.2) \times 10^{-5} \text{ cm}^{-3}$  (see Table 3), is very similar to that of the prototype GRB 090227B,  $\langle n_{\text{CBM}} \rangle = (1.90 \pm 0.20) \times 10^{-5} \text{ cm}^{-3}$ . In both the cases the CBM densities are typical of the galactic halo environment.

We turn now to the spectrum of the prompt emission using the spectral model described in Patricelli et al. (2012) with a phenomenological parameter  $\alpha = -1.11$ . From fitting the light curve in the energy range 260 keV–40 MeV, we have extended the simulation of the corresponding spectrum down to 8 keV to check overall agreement with the observed data. The final result is plotted in Figure 11, where the rebinned NaI-n6 and n9 and BGO-b1 data in the  $\Delta T_2$  time interval show their agreement with the simulation; the lower panel in Figure 11 shows the residuals of the data around the fireshell simulated spectrum.

The fireshell approach is different from the fireball one, where the sharp luminosity variations observed in the prompt emission are attributed to the prolonged and variable activity of the “inner engine” (see, e.g., Rees & Meszaros 1994; Ramirez-Ruiz & Fenimore 2000; Piran 2004).

In the fireshell model, the observed time variability of the prompt emission is produced by the interaction of the accelerated baryons of the fireshell with the CBM “clumps” (see, e.g., Ruffini et al. 2002, 2006; Patricelli et al. 2012). The issue of the time variability in GRB light curves has been long debated. Zhang et al. (2006) and Nakar & Granot (2007) indicated difficulties in producing short time variability from CBM inhomogeneities. The opposite point of view has been expressed by Dermer & Mitman (1999) and Dermer (2006, 2008). In the fireshell model it has been shown that, from the correct computation of the equations of motion of the shell, of the EQTS, and of the Lorentz factor (Bianco & Ruffini 2005a, 2005b, and Section 3), the short time scale variability of GRB light curves occurs in regimes with the larger values of the Lorentz factor, when the total visible area of the emission region is very small and “dispersion” in arrival time of the luminosity peaks is negligible. Therefore the short time scale variability indeed can be produced by the CBM inhomogeneities (see Section 3 in Patricelli et al. 2012). This has been verified in the present case of GRB 140619B, where the values of the Lorentz factor  $\Gamma$  and the total transversal size of the fireshell visible area  $d_v$  at the initial radius of the CBM cloud are explicitly indicated in Table 3. These values of  $d_v$  are smaller than the thickness of the inhomogeneities ( $\Delta r \approx 10^{16}$ – $10^{17}$  cm) and fully justify the adopted spherical symmetry approximation (Ruffini et al. 2002, 2006; Patricelli et al. 2012). Consequently, a finer description of each substructure in the spikes observed in the light curve is not necessary and does not

change the substantial agreement of the model with the observational data, which is provided by the average densities and the filling factors in Table 3.

### 6.3. The Progenitor System

In analogy with the case of GRB 090227B (see, e.g., Muccino et al. 2013; Oliveira et al. 2014), we conclude that the progenitor of GRB 140619B is a NS–NS merger. As a lower limit, we have considered the simplest case by assuming two NSs with the same mass  $M_{\text{NS}}$  such that the total mass would be larger than the NS critical mass  $M_{\text{crit}}^{\text{NS}}$ , e.g.,  $2M_{\text{NS}} \gtrsim M_{\text{crit}}^{\text{NS}}$ . This condition is clearly necessary for the formation of a BH and the consequent application of the fireshell model. It is also appropriate here to recall that only a subset of binary NSs mergers can fulfill this stringent requirement (see Figure 3). This will strongly affect the estimate of the rate of these family-2 short GRBs, when compared with the usual expected binary NS rate (see Section 9 and Conclusions).

Referring to the work of Belvedere et al. (2012) on nonrotating NSs in the global charge neutrality treatment with all the fundamental interactions taken into account properly, we have considered two NSs with mass  $M_{\text{NS}} = 1.34 M_{\odot} = 0.5M_{\text{crit}}^{\text{NS}}$  and corresponding radius  $R = 12.24$  km. As a working hypothesis we assume that in the NS merger the crustal material from both NSs contributes to the GRB baryon load, while the NS cores collapse to a BH. For each NS the crustal mass from the NL3 nuclear model is  $M_c = 3.63 \times 10^{-5} M_{\odot}$ , so the total NS merger crustal mass is  $M_{2c} = 2M_c = 7.26 \times 10^{-5} M_{\odot}$ . On the other hand, the baryonic mass engulfed by the  $e^+e^-$  plasma before transparency is  $M_B = E_{e^+e^-}^{\text{tot}} B/c^2 = (1.86 \pm 0.35) \times 10^{-6} M_{\odot}$ , so we can conclude that only a small fraction of the crustal mass contributes to the baryon load, namely  $M_B = (2.56 \pm 0.48)\% M_{2c}$ . This value is consistent with the global charge neutrality condition adopted in Belvedere et al. (2012). The usually adopted LCN condition leads instead to a crustal mass  $M_c^{\text{LCN}} \sim 0.2 M_{\odot}$  (see, e.g., Belvedere et al. 2012; Oliveira et al. 2014), which would be inconsistent with the small value of the baryon load inferred above.

## 7. ON THE GWS EMISSION AND THE DETECTABILITY OR ABSENCE THEREOF

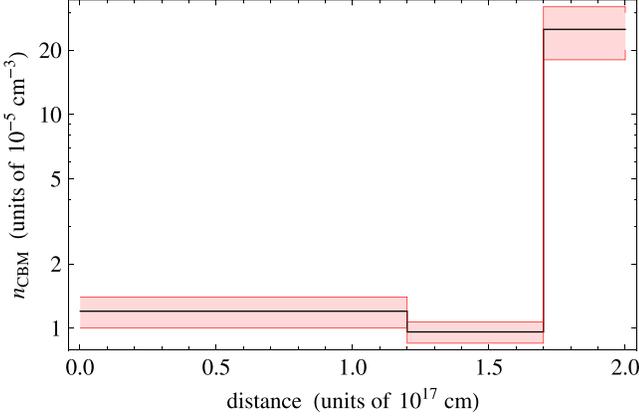
Following the previous work on GRB 090227B (Oliveira et al. 2014), we now estimate the emission of GWs of the binary NS progenitor of the short GRB 140619B using the effective-one-body (EOB) formalism (Buonanno & Damour 1999, 2000; Damour et al. 2000; Damour 2001; Damour & Nagar 2010) and assess the detectability of the emission by the Advanced LIGO interferometer.<sup>13</sup> The EOB formalism maps the conservative dynamics of a binary system of nonspinning objects onto the geodesic dynamics of a single body of reduced mass  $\mu = M_1 M_2 / M$ , with total binary mass  $M = M_1 + M_2$ . The effective metric is a modified Schwarzschild metric with a rescaled radial coordinate,  $r = c^2 \eta_2 / (GM)$ , where  $r_{12}$  is the distance between the two stars. The binary binding energy as a function of the orbital frequency  $\Omega$  is given by  $E_b(\Omega) = Mc^2 [\sqrt{1 + 2\nu(\hat{H}_{\text{eff}} - 1)} - 1]$ , where the effective Hamiltonian  $\hat{H}_{\text{eff}}^2 = A(u) + p_{\phi}^2 B(u)$  depends on the radial

<sup>13</sup> <http://www.advancedligo.mit.edu>

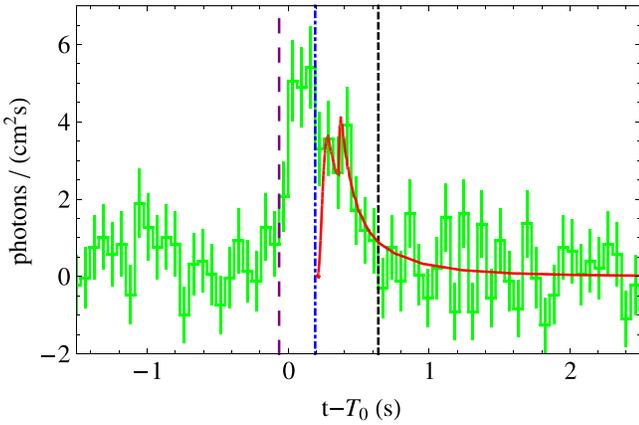
**Table 3**  
The Density and Filling Factor Masks of GRB 140619B

Cloud	Distance (cm)	$n_{\text{CBM}}$ ( $\text{cm}^{-3}$ )	$\mathcal{R}$	$\Gamma$	$d_v$ (cm)
1th	$1.50 \times 10^{15}$	$(1.2 \pm 0.2) \times 10^{-5}$	$(2.8 \pm 0.3) \times 10^{-11}$	$1.08 \times 10^4$	$2.76 \times 10^{11}$
2nd	$1.20 \times 10^{17}$	$(9.2 \pm 1.1) \times 10^{-6}$	...	$2.07 \times 10^3$	$1.16 \times 10^{14}$
3rd	$1.70 \times 10^{17}$	$(2.5 \pm 0.5) \times 10^{-4}$	$(3.5 \pm 0.6) \times 10^{-10}$	$1.84 \times 10^3$	$1.85 \times 10^{14}$

**Note.** In each column are listed, respectively, the CBM cloud, the corresponding initial radius away from the BH, the number density, the filling factor, the Lorentz factor, and the total transversal size of the fireshell visible area.

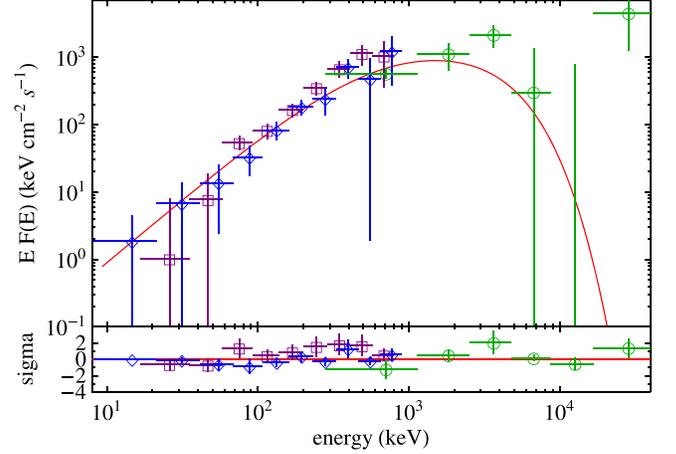


**Figure 9.** Radial CBM number density distribution of GRB 140619B (black line) and its range of validity (red shaded region).



**Figure 10.** BGO-b1 (260 keV–40 MeV) simulated light curve of the prompt emission of GRB 140619B (solid red line). Each spike corresponds to the CBM number density profile described in Table 2 and Figure 9. The blue dotted–dashed vertical line marks the end of the P-GRB emission. The purple long-dashed and the black dashed vertical lines indicate, respectively, the starting and the ending times of the  $T_{90}$  time interval. Clearly visible outside of this time interval is the background noise level. The continuation of the simulation after  $T_{90}$  is due to the residual large angle emission of the EQTS (Bianco & Ruffini 2005a, 2005b) due to the density profile indicated in Table 3.

potential  $A(u)$  of the variable  $u = 1/r$  and  $B(u) = u^2 A(u)$ , while the angular momentum for the circular orbit is given by  $p_\phi^2 = -A'(u)/[u^2 A(u)]'$ , where a prime stands for the derivative with respect to  $u$  (see, e.g., Bini & Damour 2013 for further details). In order to obtain the derivative of the effective Hamiltonian  $\hat{H}_{\text{eff}}$  as a function of  $\Omega$ , we must use the chain rule together with the relation  $\Omega = \Omega(u)$  following from the angular



**Figure 11.** Top panel: comparison between the 8–900 keV data from the NaI-n6 (purple squares) and n9 (blue diamonds) detectors, and the 260 keV–40 MeV data from the BGO-b1 detector (green circles), and the simulation within the fireshell model (solid red curve) in the time interval  $\Delta T_2$ . Bottom panel: the residuals of the above mentioned data with the simulation.

Hamilton equation of motion in the circular case  $GM\Omega(u) = (1/u)\partial H/\partial p_\phi = MA(u)p_\phi(u)u^2/(H\hat{H}_{\text{eff}})$ , where  $G$  is the gravitational constant. Finally we obtain the rate of orbital energy loss through emission of GWs from the related derivative  $dE_b/d\Omega$ .

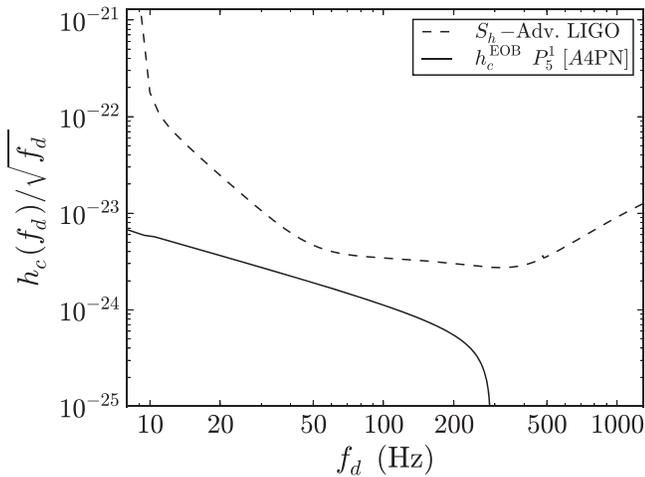
Using the well known matched filtering technique, we compute the signal-to-noise ratio (S/N) from the Fourier transform of the signal  $h(t) = h_{+}F_{+} + h_{\times}F_{\times}$ , where  $h_{+,\times}$  are functions that depend on the direction and polarization of the source and  $F_{+,\times}$  depend on the direction of the detector. By making an rms average over all possible source directions and wave polarizations, i.e.,  $\langle F_{+}^2 \rangle = \langle F_{\times}^2 \rangle = 1/5$ , we obtain (see Flanagan & Hughes 1998 for details)

$$\langle S/N^2 \rangle = \int_{f_{\min}}^{f_{\max}} df_d \frac{h_c^2(f_d)}{5f_d^2 S_h^2(f_d)}, \quad (3)$$

where  $S_h(f)$  is the strain noise spectral density (in units  $1/\sqrt{\text{Hz}}$ ) of the interferometer. We have also introduced the characteristic GW amplitude,  $h_c$ , defined using the Fourier transform of the GW form  $h(t)$ ,  $h_c(f) = f|\hat{h}(f)|$ , and it is given by

$$h_c^2(f) = \frac{2(1+z)^2}{\pi^2 d_L^2} \frac{dE_b}{df} [(1+z)f_d], \quad (4)$$

with  $z$  the cosmological redshift,  $f_d = f/(1+z)$  the GW frequency at the detector,  $f = \Omega/\pi$  the frequency in the source frame,  $f_{\min}$  the minimal bandwidth frequency of the detector, and  $f_{\max} = f_c/(1+z)$  the maximal bandwidth frequency,



**Figure 12.** Sensitivity curve of Advanced LIGO  $S_h(f)$  (dashed black curve) and the characteristic gravitational amplitude  $h_c(f_d)/\sqrt{f_d}$  (solid black curve) of the binary NS progenitor of GRB 140619B, as a function of the frequency at the detector  $f_d$ . The EOB radial potential  $A(u)$  was calculated using values for the coefficients in the 4th order post-Newtonian (PN) approximation and  $P_5^1$  is the Padé approximant of order (1, 5).

where  $f_c = \Omega_c/\pi$  is the binary contact frequency and  $d_L$  is the luminosity distance. In Figure 12 we show the strain-noise sensitivity of Advanced LIGO,  $S_h(f)$ , and the characteristic gravitational amplitude per square root frequency,  $h_c(f_d)/\sqrt{f_d}$ , both plotted as functions of the frequency at the detector  $f_d$ .

Following the above procedure we obtained for the short GRB 140619B a very low value  $\langle S/N \rangle \approx 0.21$  compared to the value  $S/N = 8$  needed for an optimal positive detection. The low value of the  $S/N$  is clearly due to the large cosmological distance to the source,  $d \approx 21$  Gpc. Although the rms-averaged  $S/N$  we have computed might improve by a factor  $\approx 5/2$  for an optimally located and polarized source (e.g.  $\langle F_+^2 \rangle = 1$  and  $\langle F_\times^2 \rangle = 0$ ) with an optimal face-on orbit ( $\cos i = 1$ ), in the case of GRB 140619B it would increase only to a maximal value  $S/N(\text{opt}) \approx 0.5$ . From the dynamics of the system, we also find that this binary emits a total energy of  $E_{\text{GW}}^T = 7.42 \times 10^{52}$  erg in gravitational radiation during the entire inspiral phase all the way up to the merger.

## 8. CONSIDERATIONS ON THE GEV EMISSION OF GRB 140619B

In addition to the analogies with GRB 090227B, GRB 140619B presents a novelty of special interest: a short-lived emission ( $\sim 4$  s) observed at energies  $\gtrsim 0.1$  GeV. The light curve of this emission shows a rising part which peaks at  $\sim 2$  s, followed by a decaying tail emission lasting another  $\sim 2$  s in the observer frame (see Figure 13(b)). Since GRB 140619B was in the LAT FoV during the entire observational period, the absence of emission after  $\sim 4$  s has been attributed to a cut-off intrinsic to the source. We divided the overall emission into four time intervals (see Figure 13(b)), each of them lasting 1 s. The corresponding spectra are best fit by power-law models. The total isotropic energy of the 0.1–100 GeV emission is  $E_{\text{LAT}} = (2.34 \pm 0.91) \times 10^{52}$  erg.

In complete analogy with the GeV emission emitted in the binary-driven hypernovae (BdHNe), we attribute this high energy radiation to the newly formed BH. This identification is clearer here in view of the absence of a supernova (SN) and the

related constant power-law emission in X-rays, when measured in the cosmological rest-frame of the BdHN (Ruffini et al. 2014, 2015; Ruffini 2015b).

The presence of this GeV emission is not a peculiarity of GRB 140619B, but is a common feature of all these family-2 short GRBs. In line with this, the apparent absence of the GeV emission in GRB 090227B has already been discussed in Section 5: it can be explained simply by the fact that this source was outside the nominal LAT FoV. The significant detection in the LLE channel and the presence of only one transient-class event with energy above 100 MeV associated with the GRB (Ackermann et al. 2013) confirms that in optimal conditions the GeV emission from GRB 090227B should have been detected.

Now consider GRB 090510, which has the characteristics of the family-2 short GRBs ( $E_{\text{iso}} > 10^{52}$  erg and  $E_{p,i} > 2$  MeV), including the presence of a high energy GeV emission lasting  $\sim 10^2$  s. This high energy emission continues up to the signal goes below the LAT threshold (Ackermann et al. 2013). The new feature of GRB 090510, among the family-2 short GRBs, is a well determined cosmological redshift inferred from the optical observations. The corresponding distance indeed coincides with the one theoretically predicted in the fireshell binary merger model (M. Muccino et al. 2015, in preparation).

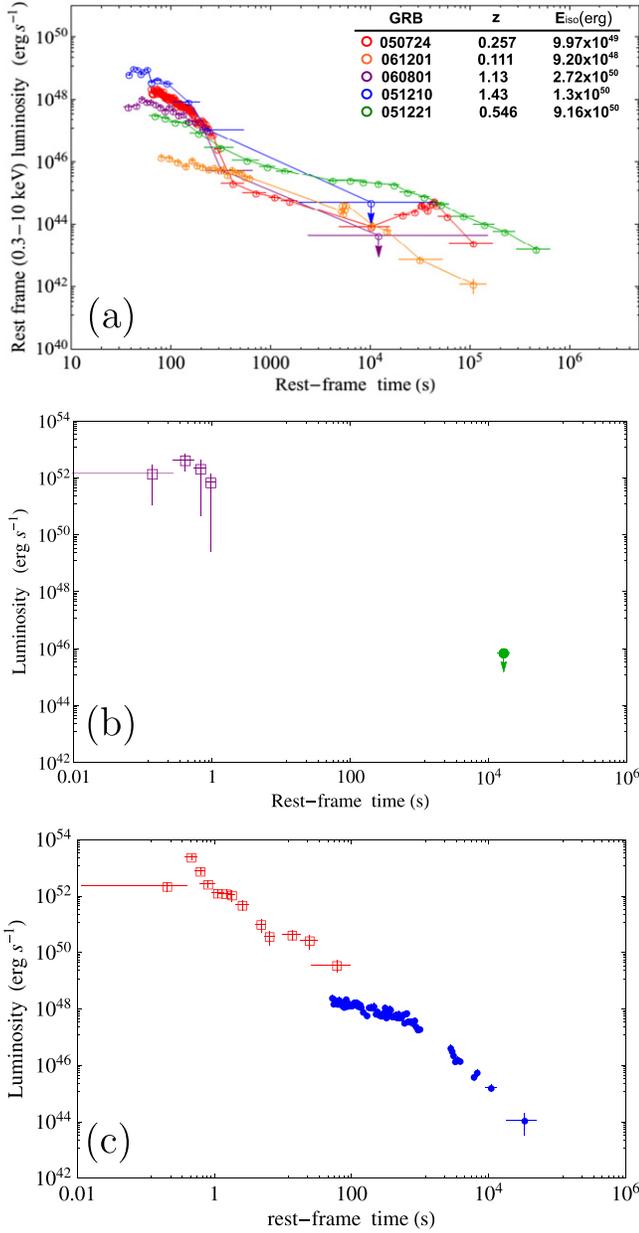
In Figure 13(a) we compare and contrast the afterglows of the traditional low energetic short GRBs (see Berger 2014, for a review) with those of the family-2 short GRB 140619B (see Figure 13(b)) and GRB 090510 (see Figure 13(c)). In Figure 1 we show the evolution of the NS–NS merger generating a family-2 short GRB. In this system the conservation laws for total energy and the total angular momentum have to be satisfied during and following the binary NS merger (J. A. Rueda et al. 2015, in preparation). One of the most important issues is the determination of the dimensionless angular momentum  $cJ/(GM^2)$  of the newly born BH (where  $J$  and  $M$  are, respectively, the BH spin angular momentum and mass). These considerations have been applied to GRB 090510 (M. Muccino et al. 2015, in preparation).

Before closing, we call attention to GRB 081024B, which we are currently addressing within the fireshell model (Y. Aimuratov et al. 2015, in preparation), and which shows all the typical features of the family-2 short GRBs, including a distinctive GeV emission. In conclusion, we can safely assert that all family-2 short GRBs, when the observational requirements are fulfilled, present a short-lived but very intense GeV emission, which in our interpretation originates from the newly formed BH.

In Table 4 we listed the redshift,  $E_{p,i}$ ,  $E_{\text{iso}}$ , and the GeV isotropic emission energy  $E_{\text{LAT}}$  in the rest-frame energy band 0.1–100 GeV of the three family-2 short GRBs discussed here. In computing  $E_{\text{iso}}$  we have inserted the energy computed in the rest-frame energy band 1–10000 keV.

## 9. THE RATE OF FAMILY-2 SHORT GRBS

With the identification of three family-2 short GRBs, namely GRB 090227B and GRB 140619B, with theoretically inferred redshifts, and GRB 090510 with a measured redshift, all of them detected by the *Fermi* satellite, we are now in a position to give an estimate of the expected rate  $\rho_0$  of such events. Following Soderberg et al. (2006) and Guetta & Della Valle (2007), for these sources we have computed the 1 s peak photon flux  $f_p$  in the energy band 1–1000 keV, which is 16.98



**Figure 13.** Top panel (a): the rebinned rest-frame 0.3–10 keV X-ray luminosities of weak short GRBs leading to massive NSs; the corresponding bursts, redshifts and energies are indicated in the legend. In their afterglows there is no regular power-law behavior at late times and no nesting (Ruffini et al. 2014). Middle panel (b): the short lived rest-frame 0.1–100 GeV isotropic luminosity light curve (purple squares) and the rest-frame 0.3–10 keV upper limit, as set from the analysis of GRB 140619B outlined in Section 3 (green circle). Bottom panel (c): the long lived rest-frame 0.1–100 GeV (red squares) and the rest-frame 0.3–10 keV (blue circles) isotropic luminosity light curves of GRB 090510.

photons  $\text{cm}^{-2} \text{s}^{-1}$  for GRB 090227B, 9.10 photons  $\text{cm}^{-2} \text{s}^{-1}$  for GRB 090510, and 4.97 photons  $\text{cm}^{-2} \text{s}^{-1}$  for GRB 140619B. From the spectral parameters for each source, we have computed  $f_p$  for various redshifts until it coincided with the corresponding threshold peak flux  $f_T$  which is the limiting peak photon flux allowing burst detection (see the analysis in Band 2003 for details). In this way we have evaluated for each source the maximum redshift  $z_{\text{max}}$  at which the burst would have been detected and, then, the corresponding maximum comoving volume  $V_{\text{max}}$ . For GRB 140619B we

**Table 4**

The Redshift, the Rest-frame Peak Spectral Energy, the Isotropic Energy  $E_{\text{iso}}$  in the Rest-frame Energy Band 1–10000 keV, and the GeV Isotropic Emission energy  $E_{\text{LAT}}$  in the Rest-frame Energy Band 0.1–100 GeV of the Four family-2 Short GRBs Discussed Here

GRB	$z$	$E_{p,i}$ (MeV)	$E_{\text{iso}}$ ( $10^{52}$ erg)	$E_{\text{LAT}}$ ( $10^{52}$ erg)
081024B	$>3.0$	$>8.2$	$>2.4$	$>2.7$
090227B	$1.61 \pm 0.14$	$5.89 \pm 0.30$	$28.3 \pm 1.5$	...
090510	$0.903 \pm 0.003$	$7.89 \pm 0.76$	$3.95 \pm 0.21$	$5.78 \pm 0.60$
140619B	$2.67 \pm 0.37$	$5.34 \pm 0.79$	$6.03 \pm 0.79$	$2.34 \pm 0.91$

**Note.** The values indicated for GRB 081024B will be Discussed in Y. Aimuravov et al. (2015, in preparation).

obtain  $f_p \equiv f_T = 1.03$  photons  $\text{cm}^{-2} \text{s}^{-1}$  at maximum redshift  $z_{140619B}^{\text{max}} = 5.49$ ; for GRB 090227B, which is the brightest one, we find  $f_p \equiv f_T = 1.68$  photons  $\text{cm}^{-2} \text{s}^{-1}$  at a maximum redshift  $z_{090227B}^{\text{max}} = 5.78$ ; finally, for GRB 090510, we get  $f_p \equiv f_T = 1.96$  photons  $\text{cm}^{-2} \text{s}^{-1}$  at a maximum redshift  $z_{090510}^{\text{max}} = 2.25$ . Correspondingly we have computed  $V_{\text{max}}$ .

The empirical rate can be evaluated as

$$\rho_0 = \left( \frac{\Omega_F}{4\pi} \right)^{-1} \frac{N}{V_{\text{max}} T_F}, \quad (5)$$

where  $N = 3$  is the number of identified energetic NS–NS short bursts,  $\Omega_F \approx 9.6$  sr is the average Fermi solid angle, and  $T = 6$  years is the Fermi observational period. We infer a local rate of  $\rho_0 = (2.6^{+4.1}_{-1.9}) \times 10^{-4} \text{ Gpc}^{-3} \text{ yr}^{-1}$ , where the attached errors are determined from the 95% confidence level of the Poisson statistic (Gehrels 1986). At  $z \geq 0.9$ , the above inferred rate provides an expected number of events  $N_{>} = 4^{+6}_{-3}$ , which is consistent with the above three observed events during the Fermi observational period. Also at  $z \leq 0.9$  our estimate  $N_{<} = 0.2^{+0.31}_{-0.14}$  is consistent with the absence of any family-2 short GRB detection.

With the inclusion of GRB 081024B, with a theoretically estimated redshift  $z > 3$  (more details will appear in Y. Aimuravov et al. 2015, in preparation), the above rate remains stable with smaller error bars, i.e.,  $\rho_0 = (2.1^{+2.8}_{-1.4}) \times 10^{-4} \text{ Gpc}^{-3} \text{ yr}^{-1}$ . This inferred rate is different from that of the long GRBs, recently estimated to be  $\rho_{\text{L-GRB}} = 1.3^{+0.7}_{-0.6} \text{ Gpc}^{-3} \text{ yr}^{-1}$  (Wanderman & Piran 2010), and also from the estimates of the family-1 short GRBs given in the literature (without a beaming correction  $\rho_{\text{short}} = 1\text{--}10 \text{ Gpc}^{-3} \text{ yr}^{-1}$ ; see e.g., Berger 2014 and Clark et al. 2014).

Such a low rate can be explained based upon the existing data of binary NSs within our Galaxy (see Section 2). From Figure 3 we notice that only a subset of them has the sum of the masses of the components larger than the critical NS mass and can collapse to a BH in their merger process. Only this subset can lead to a family-2 short GRB.

## 10. THE FAMILY-2 SHORT GRBS AND THE $E_{p,i}$ – $E_{\text{ISO}}$ RELATION FOR SHORT GRBS

Now we discuss some general considerations for the new  $E_{p,i}$ – $E_{\text{iso}}$  relation for short GRBs (Zhang et al. 2012; Calderone et al. 2015), with a power law similar to the one of the Amati

relation for long GRBs (Amati et al. 2008), but different amplitude. This yet unexplained difference discourages the use of the Amati relation as an astronomical tool. All four family-2 short GRBs satisfy this new  $E_{p,i}-E_{\text{iso}}$  relation (see the quantities listed in Table 4). We call attention to the need to investigate the physical reasons for the validity of this universal  $E_{p,i}-E_{\text{iso}}$  relation, which appears to be satisfied by family-1 short bursts, where the binary NS merger does not lead to BH formation, and also the family-2 short bursts, where BHs are formed and reveal their presence by giving rise to the short-lived but significant GeV emission.

## 11. CONCLUSIONS

In this article we have predicted the occurrence of two different kinds of short GRBs originating from binary NS mergers, based on

- (a) the analysis of GRB 090227B, the prototype of short bursts originating from a binary NS leading to BH formation (Muccino et al. 2013),
- (b) the recent progress in the determination of the mass–radius relation of NSs and the determination of their critical mass  $M_{\text{crit}}^{\text{NS}} \approx 2.67 M_{\odot}$  (Rotondo et al. 2011; Rueda et al. 2011, 2014; Belvedere et al. 2012, 2014a, 2014b), and
- (c) the recently measured mass of PSR J0348+0432,  $M = (2.01 \pm 0.04) M_{\odot}$  (Antoniadis et al. 2013), establishing an absolute lower limit on  $M_{\text{crit}}^{\text{NS}}$ , and the remarkable information gained from radio observations of binary NS systems in our own Galaxy (Zhang et al. 2011; Antoniadis 2014).

The first kind of short GRBs, which we call family-1, are the most common ones with  $E_{\text{iso}} < 10^{52}$  erg and rest-frame spectral peak energy  $E_{p,i} < 2$  MeV, originating from binary NS mergers with merged core mass smaller than  $M_{\text{crit}}^{\text{NS}}$  and leading, therefore, to a massive NS, possibly with a companion. We identify these family-1 short bursts with the ones extensively quoted in literature (see, e.g., Berger 2014 for a review).

The second kind of short GRBs, which we call family-2, are those with  $E_{\text{iso}} > 10^{52}$  erg and harder spectra with  $E_{p,i} > 2$  MeV, originating from binary NS mergers with merged core mass larger than  $M_{\text{crit}}^{\text{NS}}$ . These family-2 short bursts satisfy the necessary condition to form a BH, following the example of the prototype GRB 090227B (Muccino et al. 2013).

The application of the fireshell model (Ruffini et al. 2001a, 2001b, 2001c) to the family-2 short GRB 140619B analyzed here has allowed the determination of the physical parameters of this source: the identification of the P-GRB emission in the early  $\sim 0.2$  s of its light curve, the theoretical cosmological redshift of  $z = 2.67 \pm 0.37$  and consequently the total burst energy  $E_{e^+e^-}^{\text{tot}} = (6.03 \pm 0.79) \times 10^{52}$  erg, the baryon load  $B = (5.52 \pm 0.73) \times 10^{-5}$ , and a Lorentz  $\Gamma$  factor at transparency  $\Gamma_{\text{tr}} = (1.08 \pm 0.08) \times 10^4$ . The analysis of the prompt emission has also led to the determination of the CBM density,  $\langle n_{\text{CBM}} \rangle = (4.7 \pm 1.2) \times 10^{-5} \text{ cm}^{-3}$ , typical of the galactic halo environment, where NS–NS binaries migrate to, due to natal kicks imparted to them at birth (see, e.g., Narayan et al. 1992; Bloom et al. 1999; Fryer et al. 1999; Belczynski et al. 2006; Berger 2014), clearly supporting the binary NS merger hypothesis of this source. Unexpectedly, we have found the existence of a short-lived and

very intense GeV emission, just after the P-GRB occurrence and during and after the prompt emission phase, which has led us to conclude that this high energy emission originates from the newly formed BH.

While this article was being refereed, we have discovered three additional examples of these family-2 short bursts: GRB 081024B, GRB 090510, and GRB 090227B. These have given evidence that all these family-2 short bursts indeed show the existence of high energy emission, with the sole exception of GRB 090227B, which at the time of the observation was outside the nominal LAT FOV.

In summary we formulate some norms and theoretical predictions.

- (1) All family-1 short GRBs have an extended X-ray afterglow (see, e.g., Figure 13(a) and Berger 2014). When computed in the rest-frame 0.3–10 keV energy band they do not show any specific power-law behavior (Pisani et al. 2013) or the “nesting” properties (Ruffini et al. 2014) which have been discovered in some long GRBs. We predict that family-1 short GRBs, originating from a binary merger to a massive NS, should never exhibit high energy emission. The upper limit of  $10^{52}$  erg can be simply understood in terms of a merger leading to a massive NS.
- (2) All family-2 short GRBs have been observed not to have prominent X-ray or optical afterglows. They all have short-lived but very energetic GeV emissions (see, e.g., Figures 13(b) and (c)), when LAT data are available. The upper limit of  $10^{54}$  erg can be also simply understood in terms of a merger leading to BH formation.
- (3) The high energy emission episode in family-2 short GRBs always occurs at the end of the P-GRB emission, during and after the prompt emission phase. This fact uniquely links the high energy emission to the occurrence of the newly born BH. The prompt emission phase studied within the fireshell model has also allowed the determination of a large number of essential astrophysical parameters, both of the source (e.g.,  $E_{e^+e^-}^{\text{tot}}$  and  $B$ ) and of the CBM (e.g.,  $\alpha$ ,  $n_{\text{CBM}}$ , and  $\mathcal{R}$ ).

It is interesting that the very simplified conditions encountered in the short GRBs in the absence of a SN event, which characterize the long GRBs (Ruffini et al. 2015), have allowed definite progress in understanding some fundamental GRB properties, e.g., the correlation of high energy emission to the BH formation. They can be adapted to the case of long GRBs. The points summarized above go a long way toward reaching a better understanding of family-1 and family-2 long GRBs (Ruffini et al. 2015), as well as of the BdHNe (Ruffini et al. 2014). We are confident that GRB 140619B is one of the best examples of short GRBs obtained with the current space technology. We sincerely hope that the results of our research will lead to new missions with greater collecting area and time resolution in X- and gamma-rays.

We thank the referee for requesting additional observational support for our theoretical fireshell binary merger model. This has motivated us to improve the connection of our theoretical work with the observations, resulting in this new version of the manuscript. We are especially grateful to S. Campana and C. L. Fryer for useful suggestions in improving some conceptual and observational arguments, and to L. Izzo for the detailed analysis of the high energy emission. M.E., M.K. and F.G.O.

are supported by the Erasmus Mundus Joint Doctorate Program by grant Nos. 2012-1710, 2013-1471, and 2012-1710, respectively, from the EACEA of the European Commission. A.V.P. and E.Z. acknowledge the support by the International Cooperation Program CAPES-ICRANet financed by CAPES-Brazilian Federal Agency for Support and Evaluation of Graduate Education within the Ministry of Education of Brazil. This work made use of data supplied by the UK *Swift* Science Data Centre at the University of Leicester.

## REFERENCES

- Ackermann, M., Ajello, M., Asano, K., et al. 2013, *ApJS*, **209**, 11
- Aksenov, A. G., Ruffini, R., & Vereshchagin, G. V. 2007, *PhRvL*, **99**, 125003
- Amati, L., Guidorzi, C., Frontera, F., et al. 2008, *MNRAS*, **391**, 577
- Antoniadis, J. 2015, in *Gravitational Wave Astrophysics*, ed. C. F. Sopuerta (Cham, Switzerland: Springer International Publishing), 1
- Antoniadis, J., Freire, P. C. C., Wex, N., et al. 2013, *Sci*, **340**, 448
- Band, D., Mateson, J., Ford, L., et al. 1993, *ApJ*, **413**, 281
- Band, D. L. 2003, *ApJ*, **588**, 945
- Belczynski, K., Perna, R., Bulik, T., et al. 2006, *ApJ*, **648**, 1110
- Belvedere, R., Boshkayev, K., Rueda, J. A., & Ruffini, R. 2014a, *NuPhA*, **921**, 33
- Belvedere, R., Pugliese, D., Rueda, J. A., Ruffini, R., & Xue, S.-S. 2012, *NuPhA*, **883**, 1
- Belvedere, R., Rueda, J. A., & Ruffini, R. 2014b, *JKPS*, **65**, 897
- Berger, E. 2011, *NewAR*, **55**, 1
- Berger, E. 2014, *ARA&A*, **52**, 43
- Bianco, C. L., & Ruffini, R. 2005a, *ApJL*, **633**, L13
- Bianco, C. L., & Ruffini, R. 2005b, *ApJL*, **620**, L23
- Bini, D., & Damour, T. 2013, *PhRvD*, **87**, 121501
- Bloom, J. S., Prochaska, J. X., Pooley, D., et al. 2006, *ApJ*, **638**, 354
- Bloom, J. S., Sigurdsson, S., & Pols, O. R. 1999, *MNRAS*, **305**, 763
- Boguta, J., & Bodmer, A. R. 1977, *NuPhA*, **292**, 413
- Buonanno, A., & Damour, T. 1999, *PhRvD*, **59**, 084006
- Buonanno, A., & Damour, T. 2000, *PhRvD*, **62**, 064015
- Burrows, D., Hill, J., Nousek, J., et al. 2005, *SSRv*, **120**, 165
- Calderone, G., Ghirlanda, G., Ghisellini, G., et al. 2015, *MNRAS*, **448**, 403
- Cherubini, C., Geralico, A., J. A. Rueda, H., & Ruffini, R. 2009, *PhRvD*, **79**, 124002
- Cipolletta, F., Cherubini, C., Filippi, S., Rueda, J. A., & Ruffini, R. 2015, *PhRvD*, submitted (arXiv:1506.05926)
- Clark, J., Evans, H., Fairhurst, S., et al. 2014, arXiv:1409.8149
- Connaughton, V., Zhang, B.-B., Fitzpatrick, G., & Roberts, O. 2014, GCN, 16419, 1
- Damour, T. 2001, *PhRvD*, **64**, 124013
- Damour, T., Jaranowski, P., & Schäfer, G. 2000, *PhRvD*, **62**, 084011
- Damour, T., & Nagar, A. 2010, *PhRvD*, **81**, 084016
- Damour, T., & Ruffini, R. 1975, *PhRvL*, **35**, 463
- Deneva, J. S., Freire, P. C. C., Cordes, J. M., et al. 2012, *ApJ*, **757**, 89
- Dermer, C. D. 2006, *NCimB*, **121**, 1331
- Dermer, C. D. 2008, *ApJ*, **684**, 430
- Dermer, C. D., & Mitman, K. E. 1999, *ApJL*, **513**, L5
- Dezalay, J.-P., Barat, C., Talon, R., et al. 1992, AIP Conf. Ser. 265, ed. W. S. Paciesas & G. J. Fishman (Melville, NY: AIP), 304
- Eichler, D., Livio, M., Piran, T., & Schramm, D. N. 1989, *Natur*, **340**, 126
- Flanagan, É É., & Hughes, S. A. 1998, *PhRvD*, **57**, 4535
- Fong, W., Berger, E., & Fox, D. B. 2010, *ApJ*, **708**, 9
- Fryer, C. L., Belczynski, K., Ramirez-Ruiz, E., et al. 2015, *ApJ*, submitted (arXiv:1504.07605)
- Fryer, C. L., Rueda, J. A., & Ruffini, R. 2014, *ApJL*, **793**, L36
- Fryer, C. L., Woosley, S. E., & Hartmann, D. H. 1999, *ApJ*, **526**, 152
- Gehrels, N. 1986, *ApJ*, **303**, 336
- Goodman, J. 1986, *ApJL*, **308**, L47
- Guetta, D., & Della Valle, M. 2007, *ApJL*, **657**, L73
- Haensel, P., Potekhin, A. Y., & Yakovlev, D. G. (ed.) 2007, *Astrophysics and Space Science Library 326, Neutron Stars 1: Equation of State and Structure* (New York: Springer)
- Han, W.-B., Ruffini, R., & Xue, S.-S. 2012, *PhRvD*, **86**, 084004
- Iwakiri, W., Tashiro, M., Terada, Y., et al. 2014, GCN, 16457, 1
- Izzo, L., Bernardini, M. G., Bianco, C. L., et al. 2010, *JKPS*, **57–3**, 551
- Izzo, L., Rueda, J. A., & Ruffini, R. 2012a, *A&A*, **548**, L5
- Izzo, L., Ruffini, R., Penacchioni, A. V., et al. 2012b, *A&A*, **543**, A10
- Klebesadel, R. W. 1992, in *Gamma-Ray Bursts—Observations, Analyses and Theories*, ed. C. Ho, R. I. Epstein, & E. E. Fenimore (Cambridge: Cambridge Univ. Press)
- Kocevski, D., Longo, F., Vianello, G., et al. 2014, GCN, 16421, 1
- Kopač, D., D’Avanzo, P., Melandri, A., et al. 2012, *MNRAS*, **424**, 2392
- Kouveliotou, C., Meegan, C. A., Fishman, G. J., et al. 1993, *ApJL*, **413**, L101
- Kramer, M. 2014, *IJMPD*, **23**, 30004
- Lattimer, J. M. 2012, *ARNPS*, **62**, 485
- Lawrence, S., Tervala, J. G., Bedaque, P. F., & Miller, M. C. 2015, *ApJ*, in press (arXiv:1505.00231)
- Lee, W. H., Ramirez-Ruiz, E., & Page, D. 2004, *ApJL*, **608**, L5
- Maselli, A., & D’Avanzo, P. 2014, GCN, 16424, 1
- Meegan, C., Lichti, G., Bhat, P. N., et al. 2009, *ApJ*, **702**, 791
- Meszáros, P. 2006, *RPPH*, **69**, 2259
- Meszáros, P., & Rees, M. J. 1997, *ApJL*, **482**, L29
- Muccino, M., Ruffini, R., Bianco, C. L., Izzo, L., & Penacchioni, A. V. 2013, *ApJ*, **763**, 125
- Nakar, E., & Granot, J. 2007, *MNRAS*, **380**, 1744
- Narayan, R., Paczynski, B., & Piran, T. 1992, *ApJL*, **395**, L83
- Narayan, R., Piran, T., & Shemi, A. 1991, *ApJL*, **379**, L17
- Oliveira, F. G., Rueda, J. A., & Ruffini, R. 2014, *ApJ*, **787**, 150
- Paczynski, B. 1986, *ApJL*, **308**, L43
- Patricelli, B., Bernardini, M. G., Bianco, C. L., et al. 2012, *ApJ*, **756**, 16
- Piran, T. 2004, *RvMP*, **76**, 1143
- Piran, T. 2005, *RvMP*, **76**, 1143
- Pisani, G. B., Izzo, L., Ruffini, R., et al. 2013, *A&A*, **552**, L5
- Preparata, G., Ruffini, R., & Xue, S. 1998, *A&A*, **338**, L87
- Ramirez-Ruiz, E., & Fenimore, E. E. 2000, *ApJ*, **539**, 712
- Rees, M. J., & Meszáros, P. 1994, *ApJL*, **430**, L93
- Rhoades, C. E., & Ruffini, R. 1974, *PhRvL*, **32**, 324
- Rosswog, S., Ramirez-Ruiz, E., & Davies, M. B. 2003, *MNRAS*, **345**, 1077
- Rotondo, M., Rueda, J. A., Ruffini, R., & Xue, S.-S. 2011, *PhLB*, **701**, 667
- Rueda, J. A., & Ruffini, R. 2012, *ApJL*, **758**, L7
- Rueda, J. A., Ruffini, R., Wu, Y.-B., & Xue, S.-S. 2014, *PhRvC*, **89**, 035804
- Rueda, J. A., Ruffini, R., & Xue, S.-S. 2011, *NuPhA*, **872**, 286
- Ruffini, R. 2009, in *The Kerr Spacetime*, ed. D. L. Wiltshire, M. Visser, & S. Scott (Cambridge: Cambridge Univ. Press)
- Ruffini, R. 2015a, ARep, in press
- Ruffini, R. 2015b, Proc. of the 2nd Cesar Lattes Meeting, ed. U. Barres de Almeida et al., in press
- Ruffini, R., Aksenov, A. G., Bernardini, M. G., et al. 2009, AIP Conf. Ser. 1132, ed. M. Novello & S. Perez (Melville, NY: AIP), 199
- Ruffini, R., Bernardini, M. G., Bianco, C. L., et al. 2005, AIP Conf. Ser. 782, Xlth Brazilian School of Cosmology and Gravitation, ed. M. Novello & S. E. Perez Bergliaffa (Melville, NY: AIP), 42
- Ruffini, R., Bernardini, M. G., Bianco, C. L., et al. 2006, AIP Conf. Ser. 836, Gamma-Ray Bursts in the Swift Era, ed. S. S. Holt, N. Gehrels & J. A. Nousek (Melville, NY: AIP), 103
- Ruffini, R., Bianco, C. L., Chardonnet, P., Frascchetti, F., & Xue, S. 2002, *ApJL*, **581**, L19
- Ruffini, R., Bianco, C. L., Frascchetti, F., Xue, S.-S., & Chardonnet, P. 2001a, *ApJL*, **555**, L117
- Ruffini, R., Bianco, C. L., Frascchetti, F., Xue, S.-S., & Chardonnet, P. 2001b, *ApJL*, **555**, L113
- Ruffini, R., Bianco, C. L., Frascchetti, F., Xue, S.-S., & Chardonnet, P. 2001c, *ApJL*, **555**, L107
- Ruffini, R., Bianco, C. L., Xue, S.-S., et al. 2004, *IJMPD*, **13**, 843
- Ruffini, R., Izzo, L., Muccino, M., et al. 2013, ARep, submitted (arXiv:1311.7432)
- Ruffini, R., Muccino, M., Bianco, C. L., et al. 2014, *A&A*, **565**, L10
- Ruffini, R., Salmonson, J. D., Wilson, J. R., & Xue, S. 2000, *A&A*, **359**, 855
- Ruffini, R., Salmonson, J. D., Wilson, J. R., & Xue, S.-S. 1999, *A&A*, **350**, 334
- Ruffini, R., Vereshchagin, G., & Xue, S.-S. 2010, *PhR*, **487**, 1
- Ruffini, R., Wang, Y., Enderli, M., et al. 2015, *ApJ*, **798**, 10
- Ruffini, R., & Xue, S.-S. 2008, AIP Conf. Ser. 1059, ed. D.-S. Lee & W. Lee (Melville, NY: AIP), 72
- Ruffini, R., & Xue, S.-S. 2013, *PhLA*, **377**, 2450
- Sahu, K. C., Livio, M., Petro, L., et al. 1997, *Natur*, **387**, 476
- Salmonson, J. D., & Wilson, J. R. 2002, *ApJ*, **578**, 310
- Sari, R., Piran, T., & Narayan, R. 1998, *ApJL*, **497**, L17
- Schaefer, B. E. 2007, *ApJ*, **660**, 16
- Soderberg, A. M., Kulkarni, S. R., Nakar, E., et al. 2006, *Natur*, **442**, 1014
- Tovani, M. 1998, *ApJL*, **497**, L21
- Troja, E., King, A. R., O’Brien, P. T., Lyons, N., & Cusumano, G. 2008, *MNRAS*, **385**, L10
- van Paradijs, J., Groot, P. J., Galama, T., et al. 1997, *Natur*, **386**, 686
- Wanderman, D., & Piran, T. 2010, *MNRAS*, **406**, 1944
- Zhang, B., Fan, Y. Z., Dyks, J., et al. 2006, *ApJ*, **642**, 354
- Zhang, C. M., Wang, J., Zhao, Y. H., et al. 2011, *A&A*, **527**, A83
- Zhang, F.-W., Shao, L., Yan, J.-Z., & Wei, D.-M. 2012, *ApJ*, **750**, 88